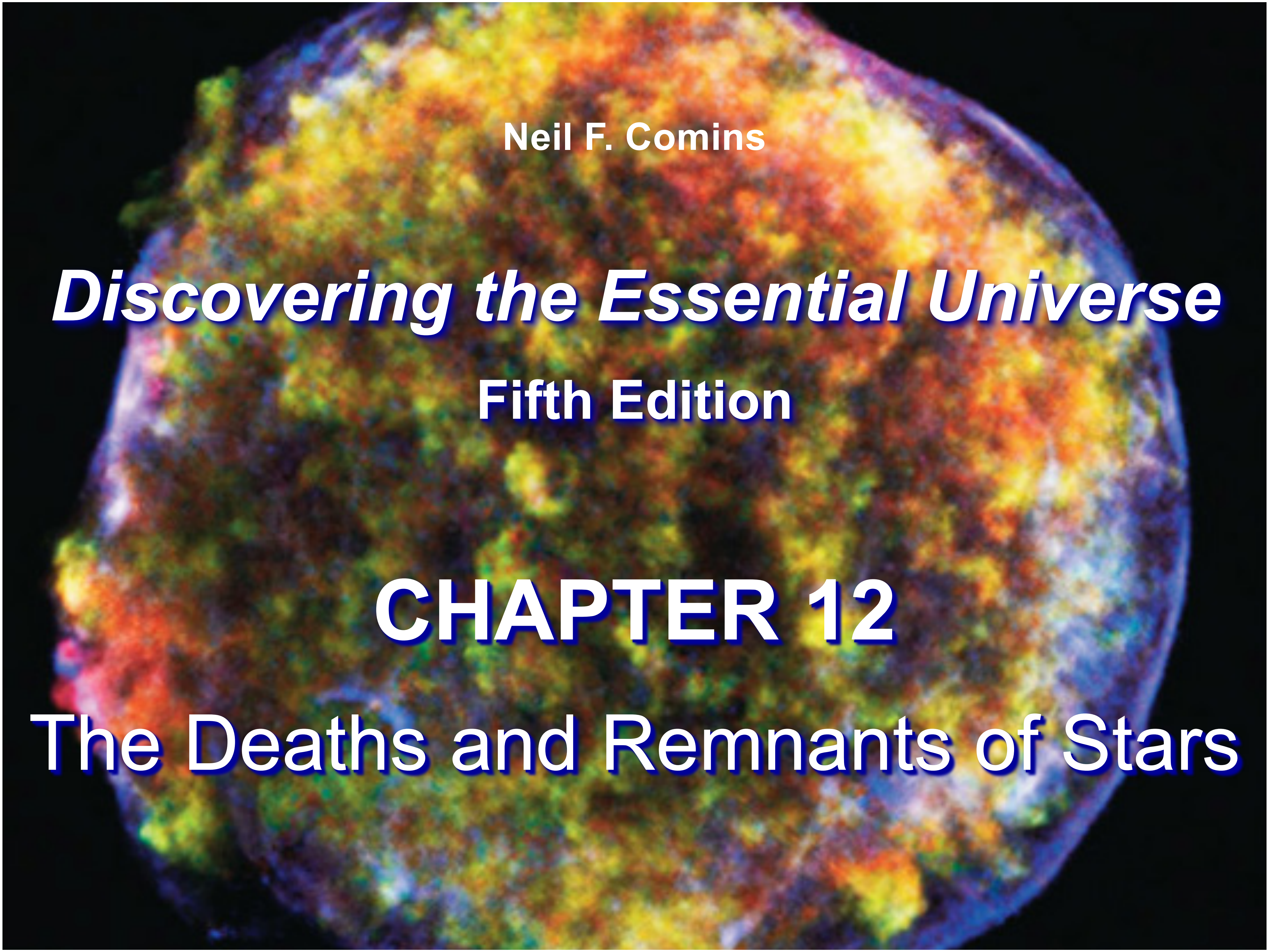


Neil F. Comins

Discovering the Essential Universe

Fifth Edition

CHAPTER 12

The Deaths and Remnants of Stars

Evolution of low-mass stars M=0.08 to 0.4 M_{sol}

- Called red dwarfs
- All H is fused into He
- Convection
- Life time > trillion years
- Most common type of stars (85%)

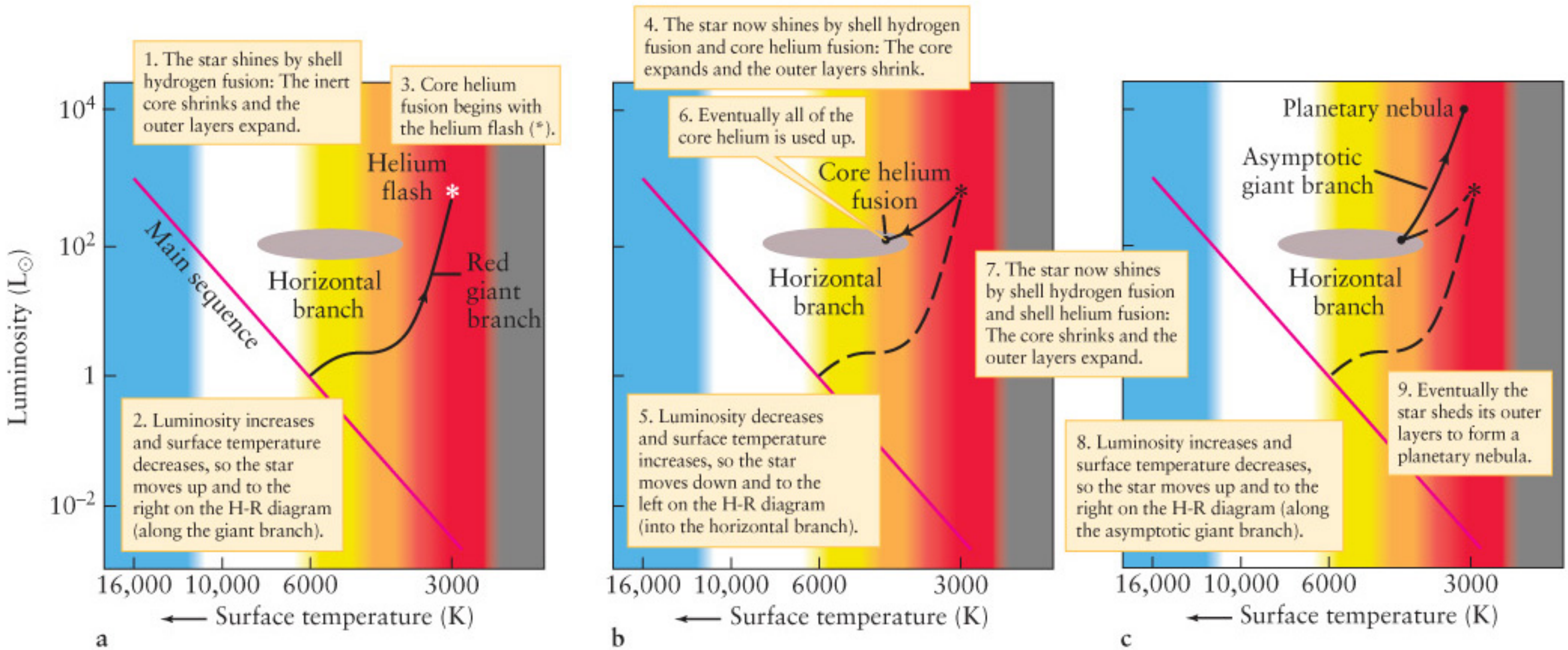
Intermediate mass stars M=0.4 to 8 M_{sol}

- Our sun is an intermediate mass star
- 4H➔ He +γ+v
- 3He➔ C +γ,
- C+He➔O +γ

High-mass stars M=8 to ~50 M_{sol}

- Fusion till Fe is produced
- Core contraction, bounce, explosion as a supernova

Post–Main-Sequence Evolution of Intermediate-Mass Stars



(a) A typical evolutionary track on the H-R diagram as a star makes the transition from the main sequence to the giant phase. The asterisk (*) shows the helium flash occurring in a low-mass star. (b) After the helium flash, the star converts its helium core into carbon and oxygen. While doing so, its core re-expands, decreasing shell fusion. As a result, the star's outer layers recontract. (c) After the helium core is completely transformed into carbon and oxygen, the core recollapses, and the outer layers re-expand, powered up the asymptotic giant branch by hydrogen shell fusion and helium shell fusion.

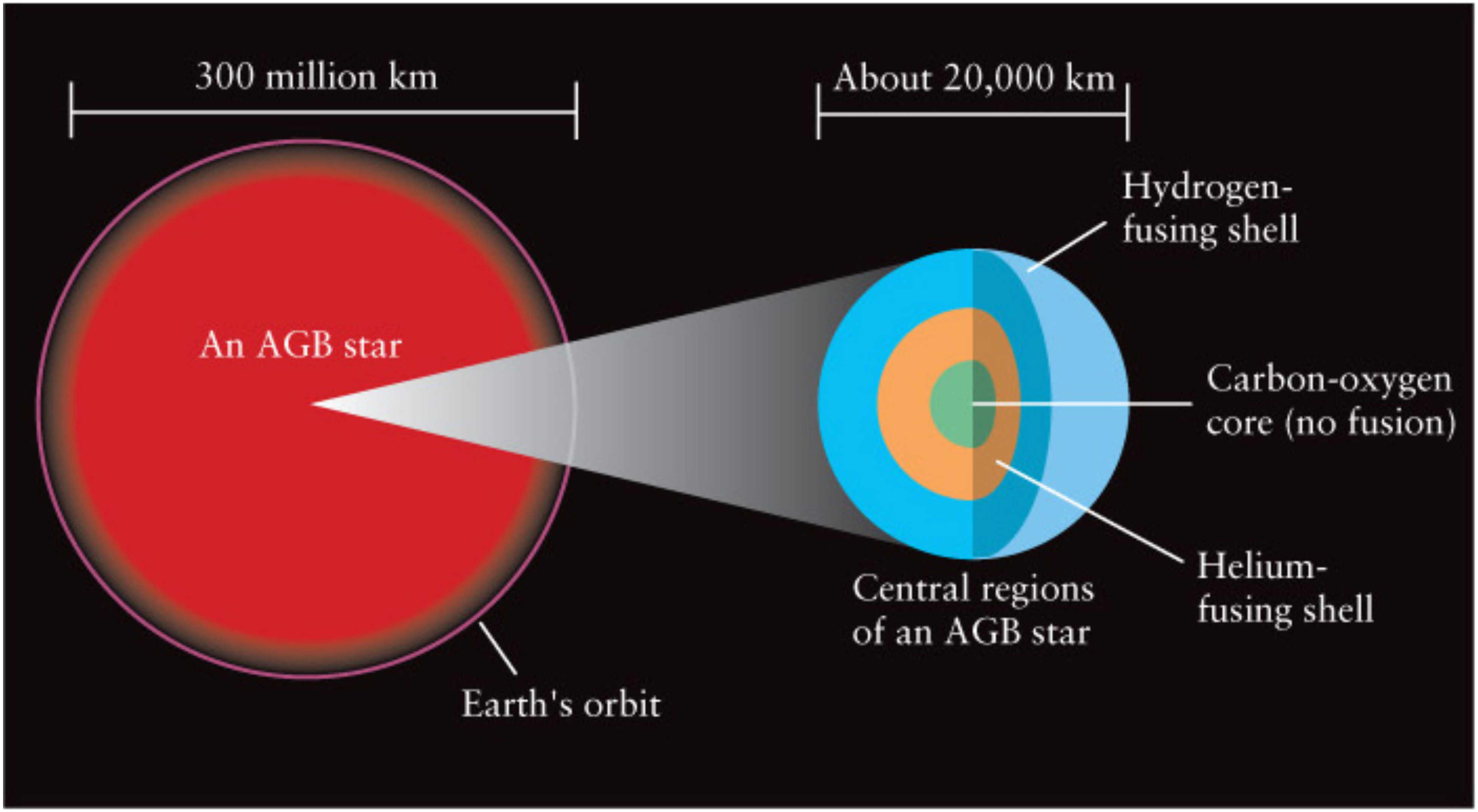
Electron degeneracy pressure

- The He-rich core of a low-mass giant is supported by electron degeneracy pressure. It is based on the:
- Pauli exclusion principle: Two identical particles cannot exist at the same place at the same time
- Electron degenerate pressure does not change with temperature

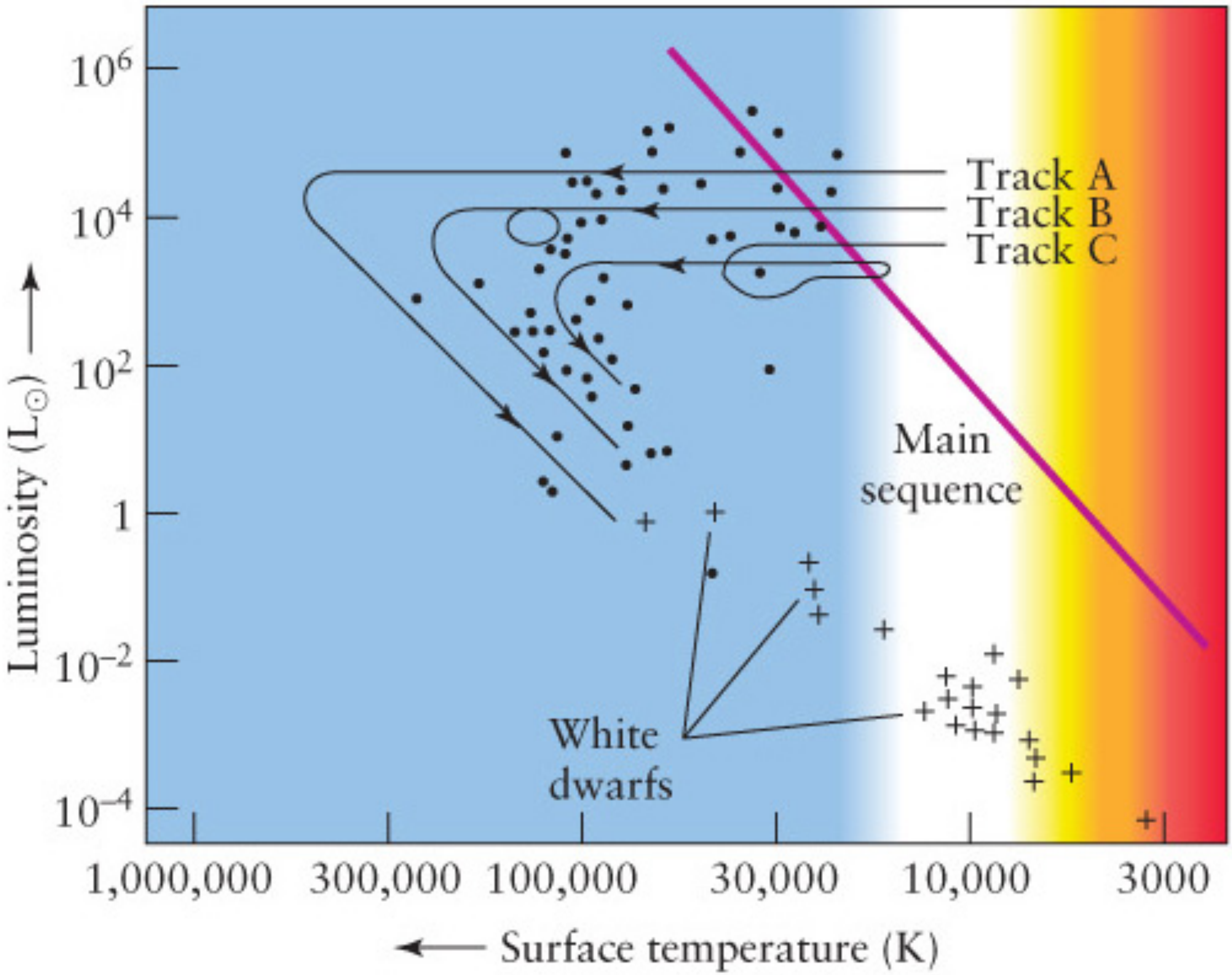
Helium flash

- Helium flash: He fusion at 100 Mill K.
- T increases but pressure constant.
- Fusion reactions on runaway for few hours.
- Luminosity increases enormously.
- When 350 Mill K is reached, He not degenerate anymore. Normal safety valve in place again.

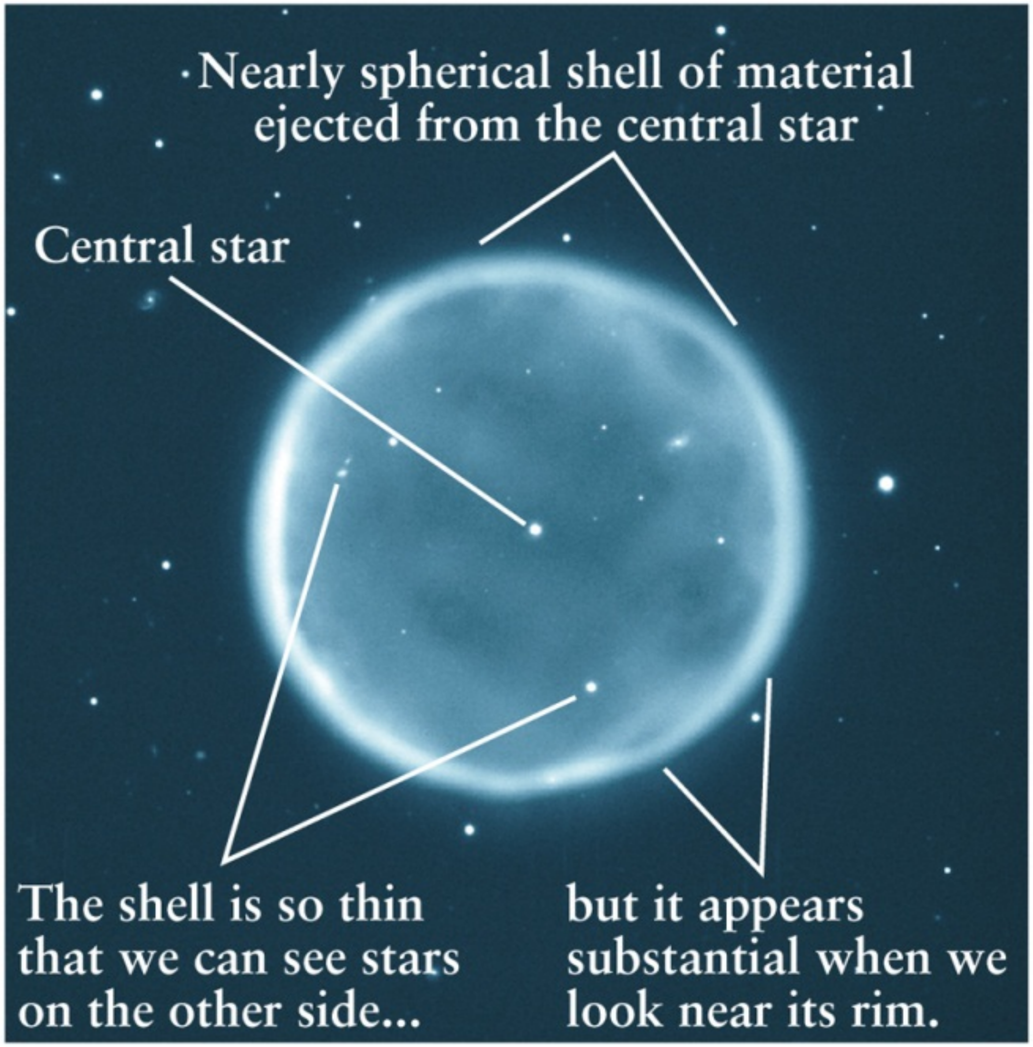
The Structure of an Old Intermediate-Mass Star

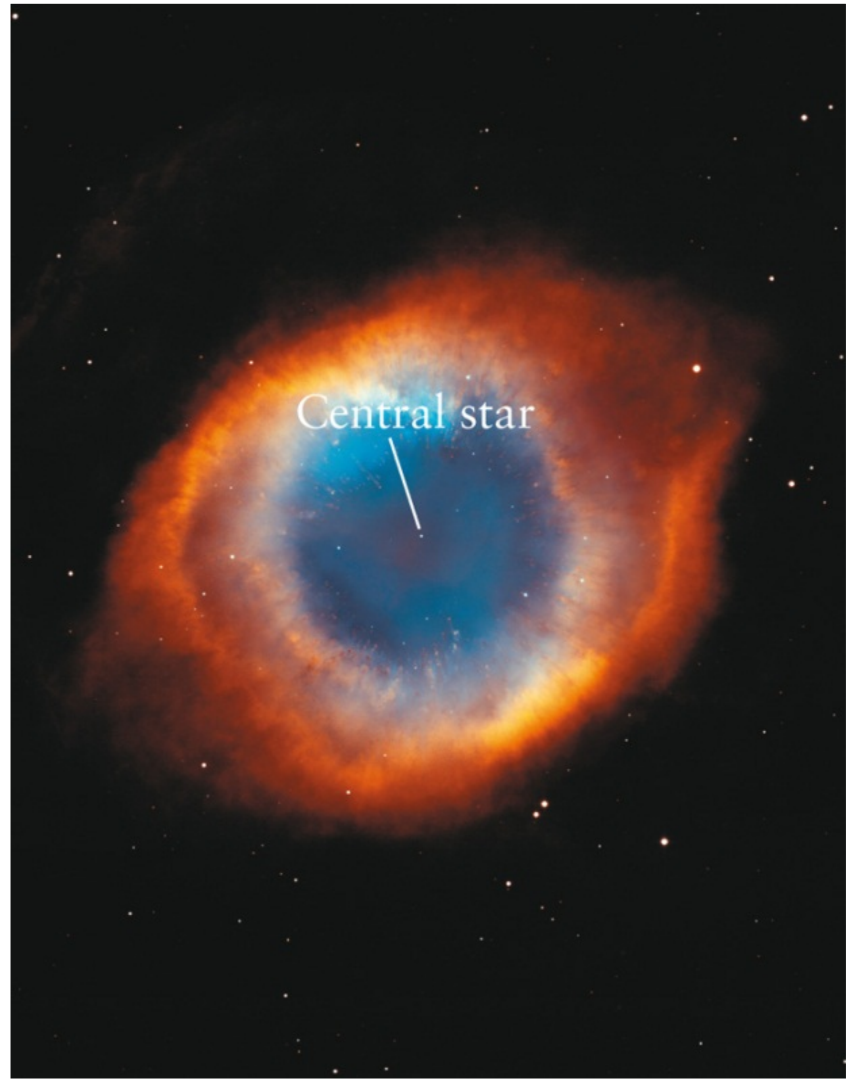


Evolution from Supergiants to White Dwarfs



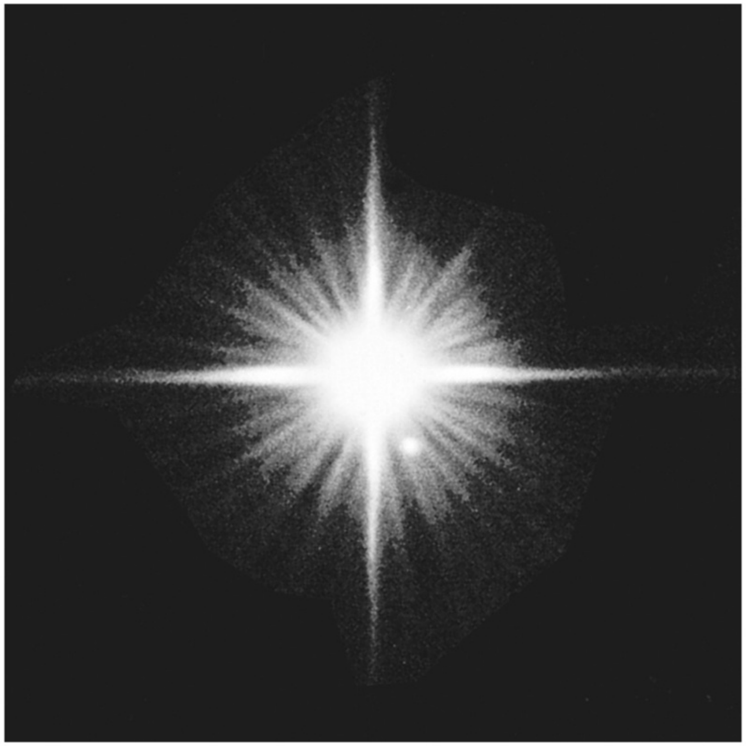
Some Shapes of Planetary Nebulae



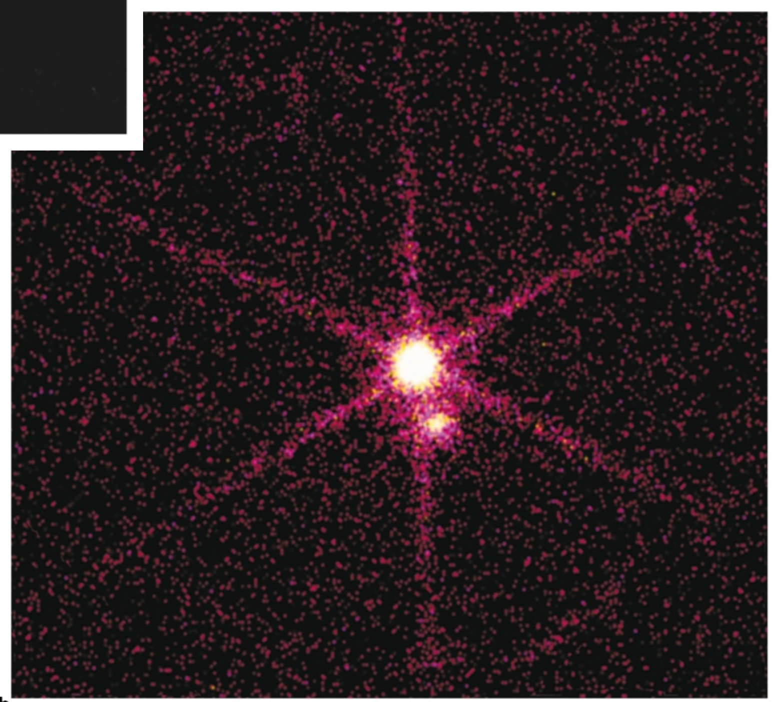


Sirius and Its White Dwarf Companion

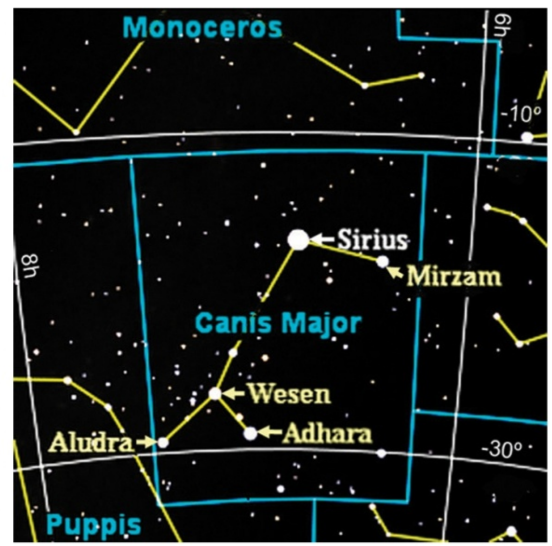
Since Sirius A (11,000 K) and Sirius B (30,000 K) are hot blackbodies, they are strong emitters of X rays.



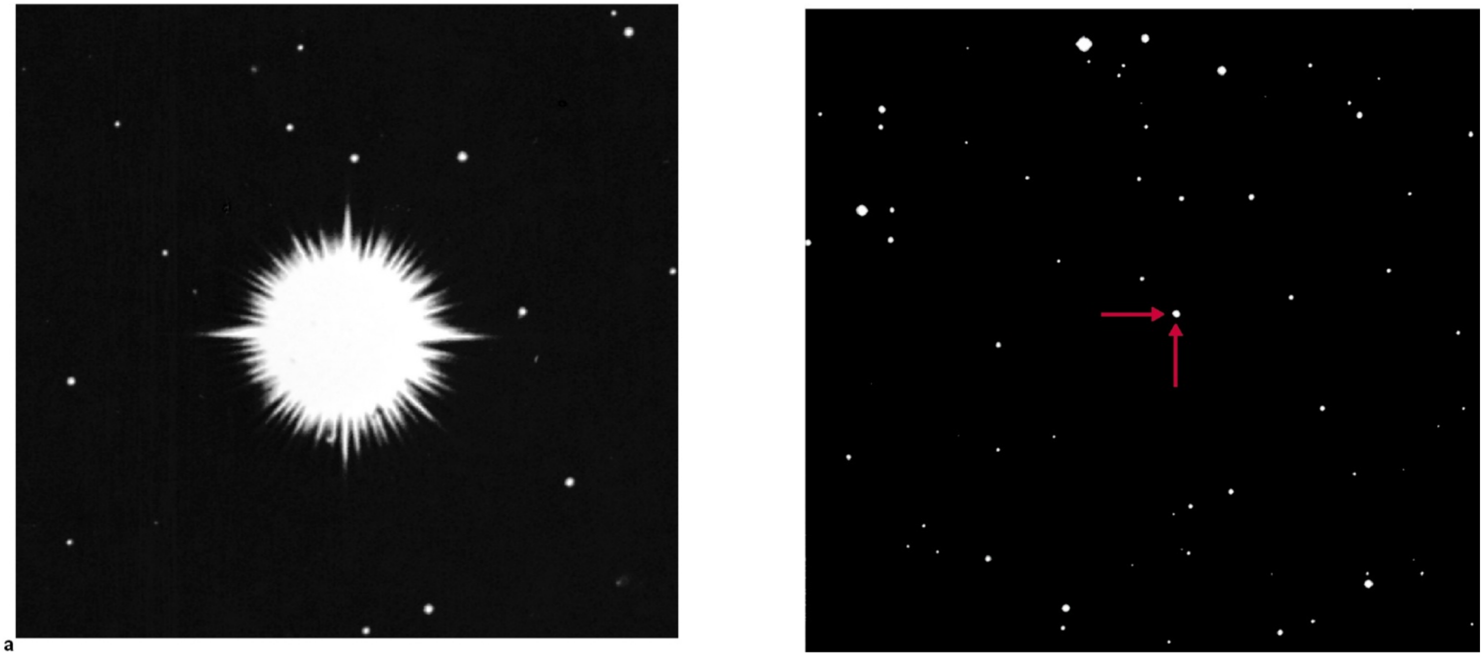
a



b



Nova Herculis 1934



a

These two pictures show a nova (a) shortly after peak brightness as a magnitude –3 star and (b) 2 months later, when it had faded to magnitude +12. Novae are named after the constellation and year in which they appear.

Two principally different kinds of SNe

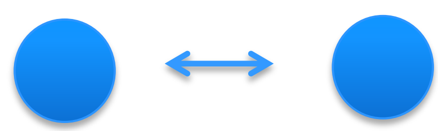
- **Type Ia supernovae**

Thermonuclear detonation of a white dwarf (WD) in a binary system

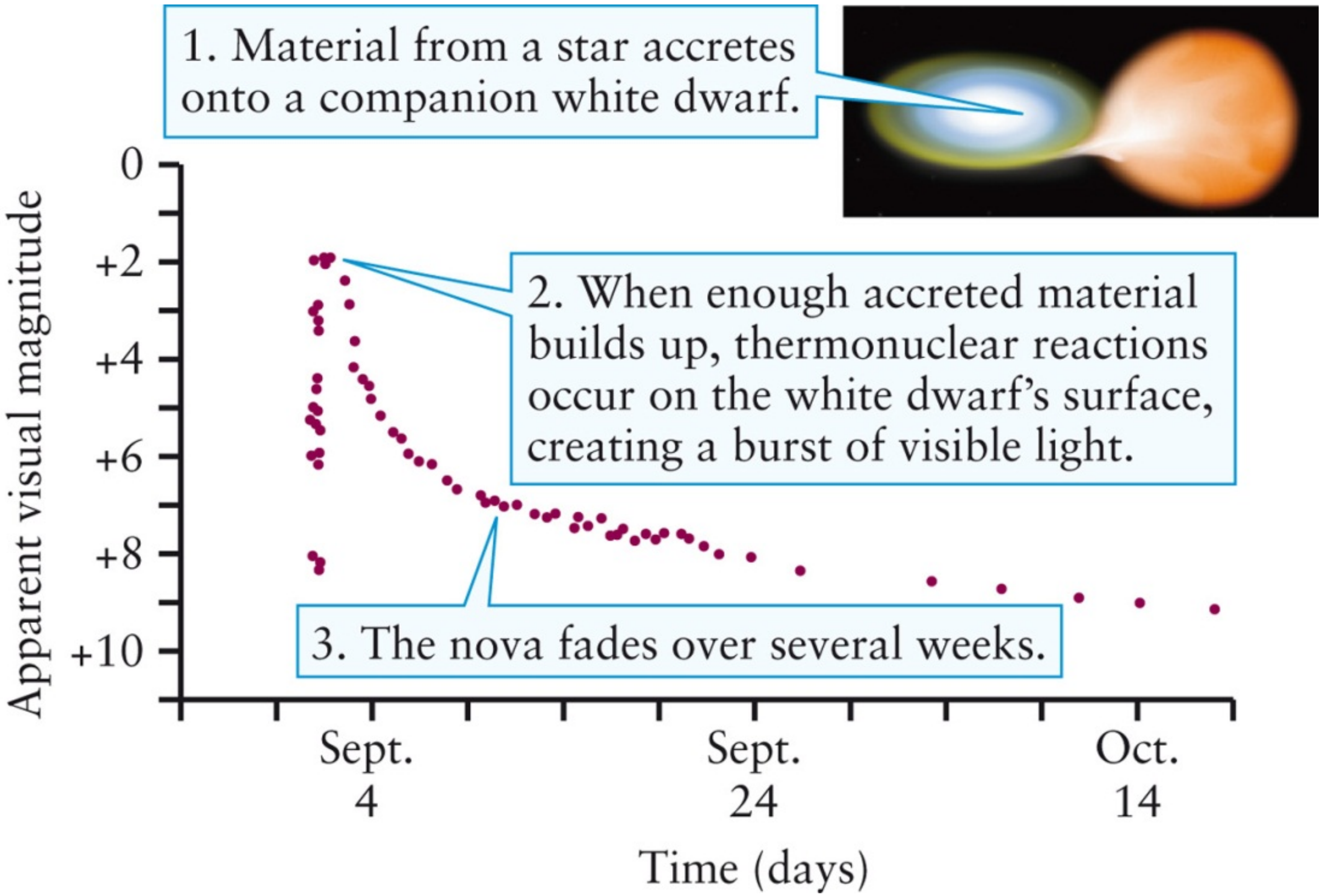
1) WD accretes material from sun-like or giant companion, reaches Chandrasek M_{sol} and detonates

/n1002/18darkenergy/accretion.jpg

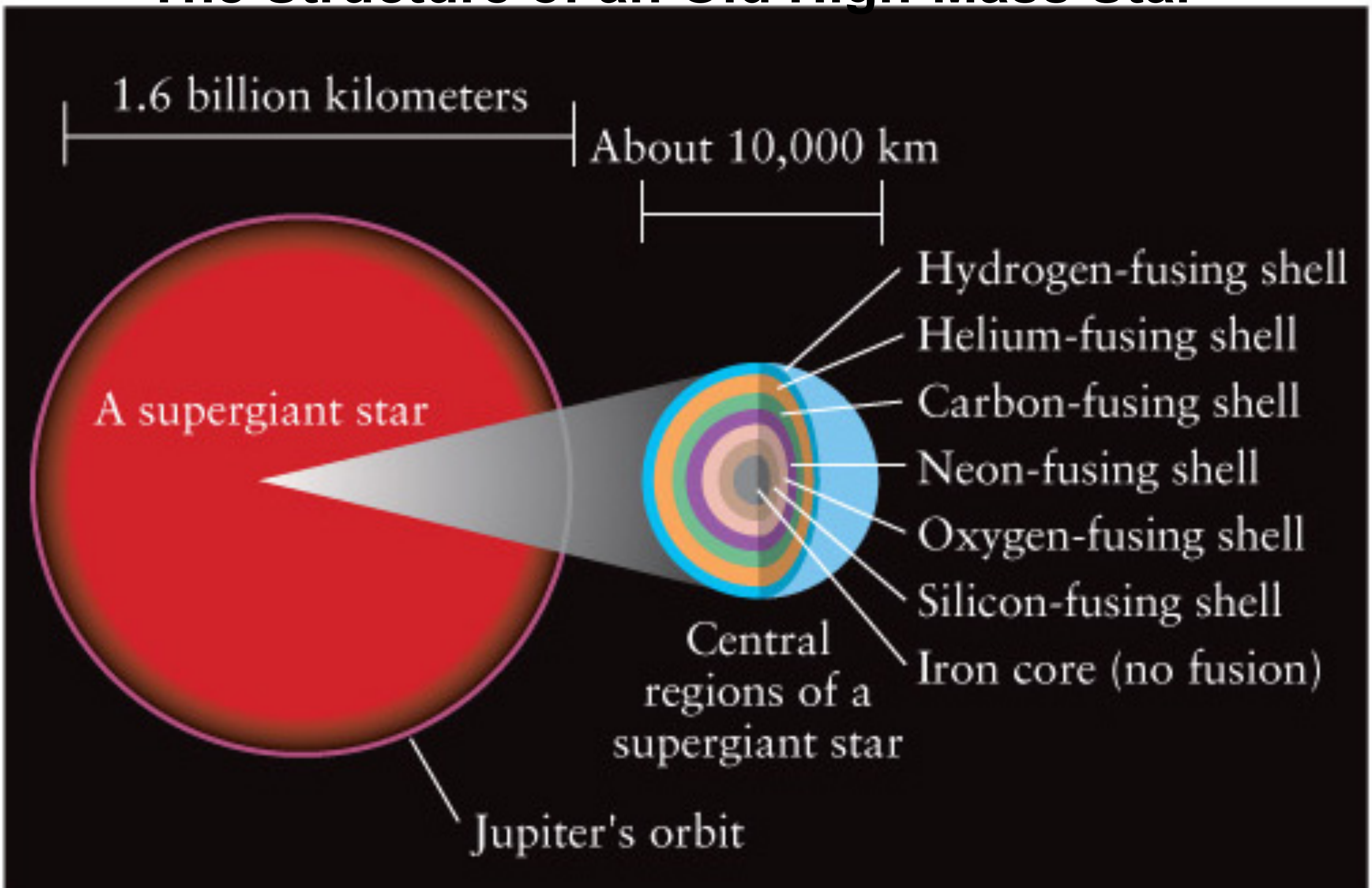
2) WD cannibalizes second WD and detonates



The Light Curve of a Nova



The Structure of an Old High-Mass Star



Near the end of its life, a high-mass star becomes a supergiant

2) Type II supernovae and other types (core collapse of a massive star)

- Core collapse of a massive star

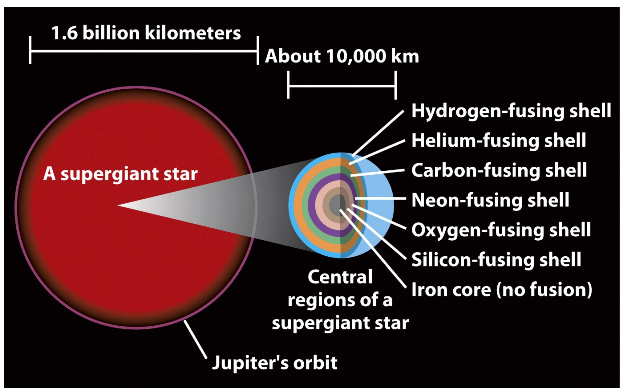
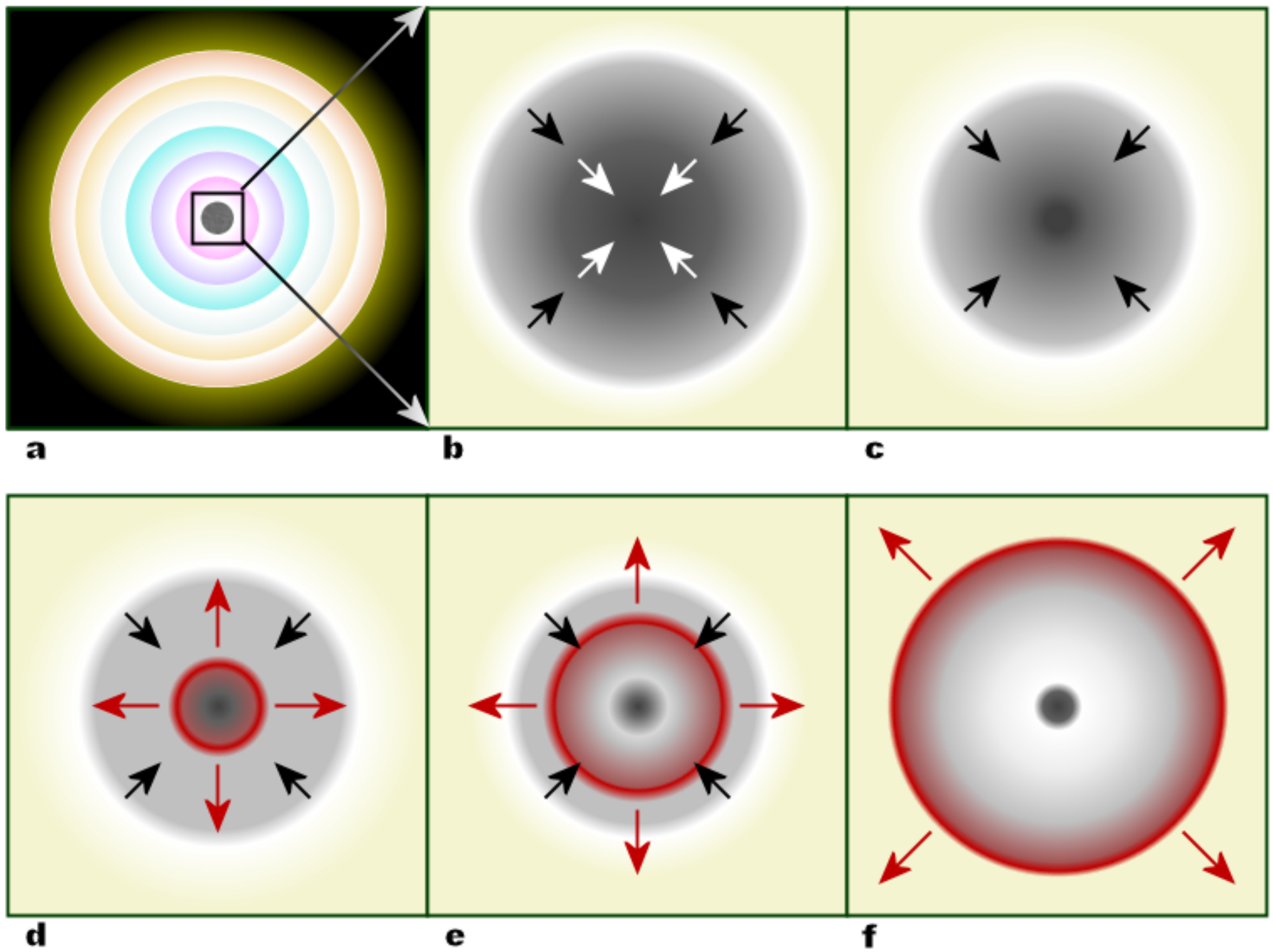
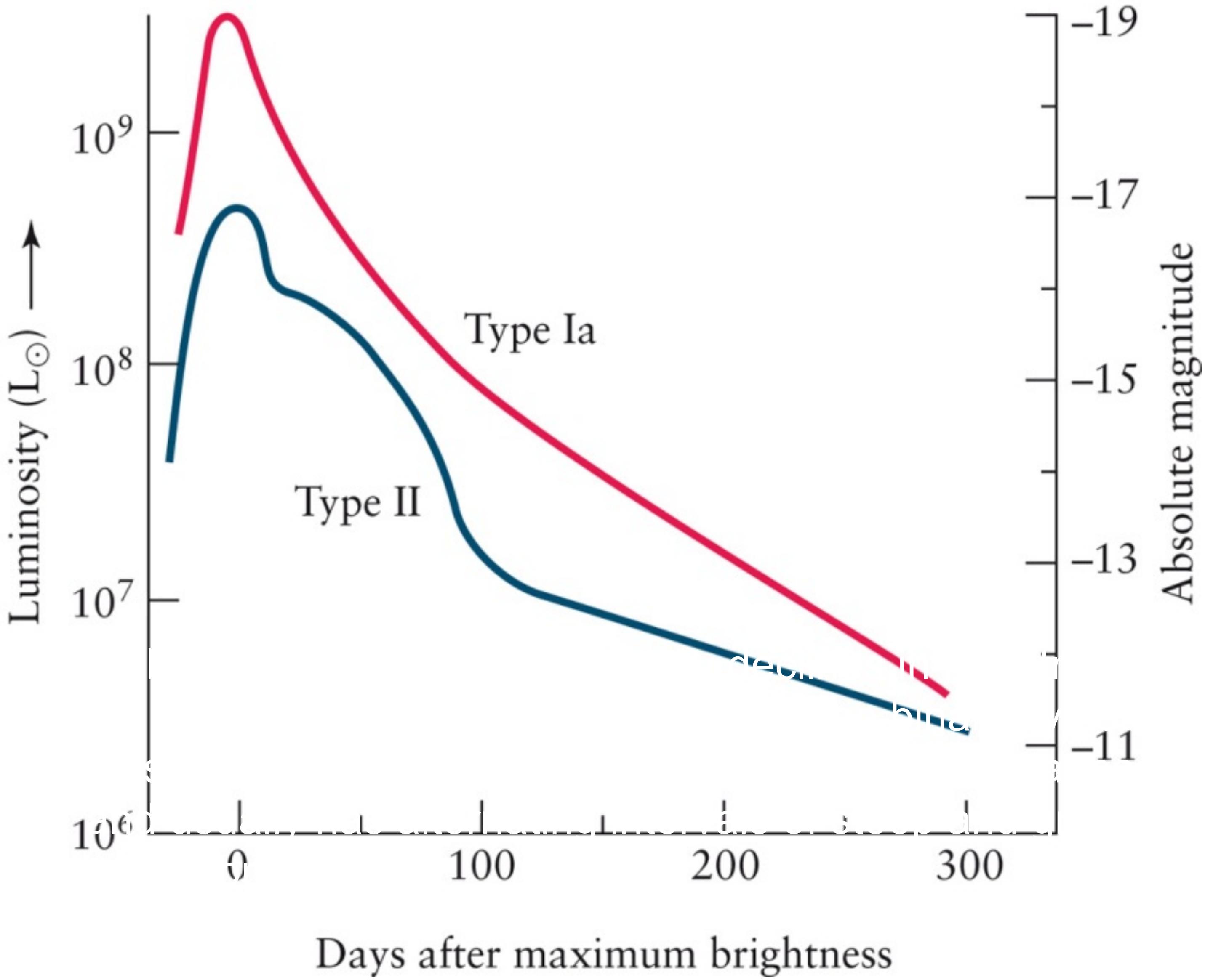


Figure 13-10
Discovering the Universe, Eighth Edition



Supernova Light Curves

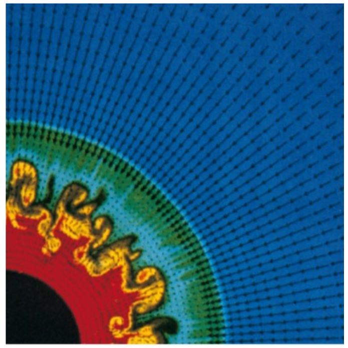


Mass Loss by a Supermassive Star

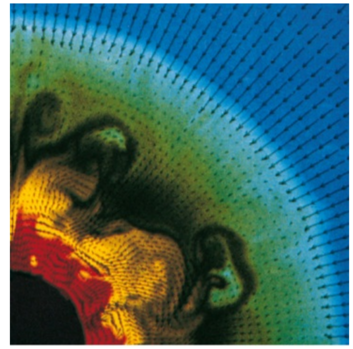
TABLE 11-1 Evolutionary Stages of a 25-$M_\odot$ Star

Stage	Central temperature (K)	Central density (kg/m^3)	Duration of stage
Hydrogen fusion	4×10^7	5×10^3	7×10^6 years
Helium fusion	2×10^8	7×10^5	5×10^5 years
Carbon fusion	6×10^8	2×10^8	600 years
Neon fusion	1.2×10^9	4×10^9	1 year
Oxygen fusion	1.5×10^9	1×10^{10}	6 months
Silicon fusion	2.7×10^9	3×10^{10}	1 day
Core collapse	5.4×10^9	3×10^{12}	0.2 second
Core bounce	2.3×10^{10}	4×10^{17}	milliseconds
Supernova explosion	about 10^9	varies	hours

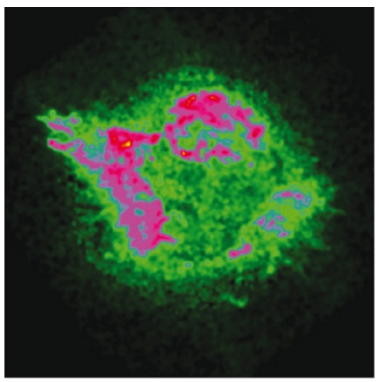
Supernovae Proceed Irregularly



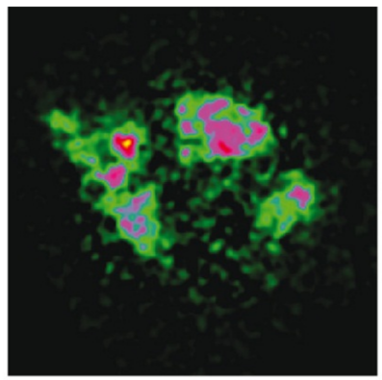
a 10 milliseconds



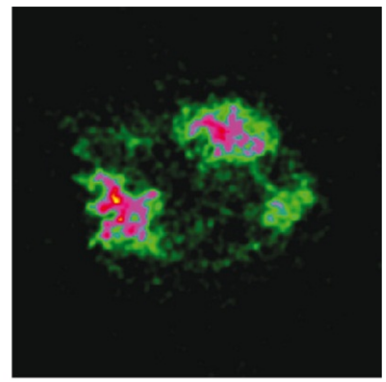
b 20 milliseconds



c Silicon

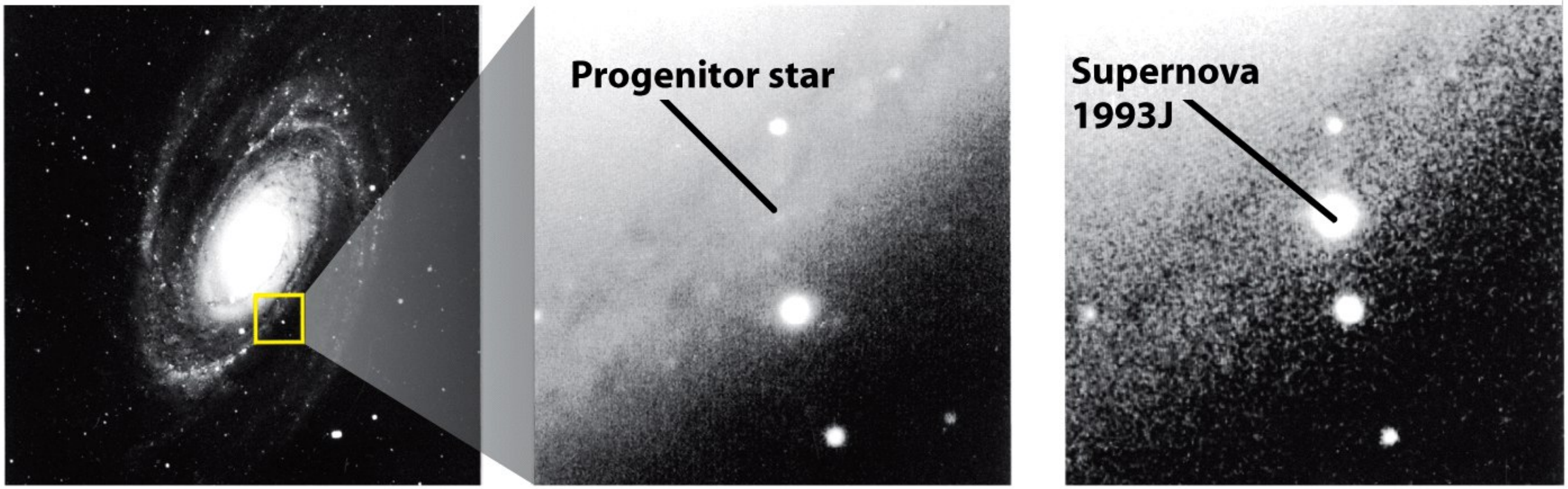


d Calcium



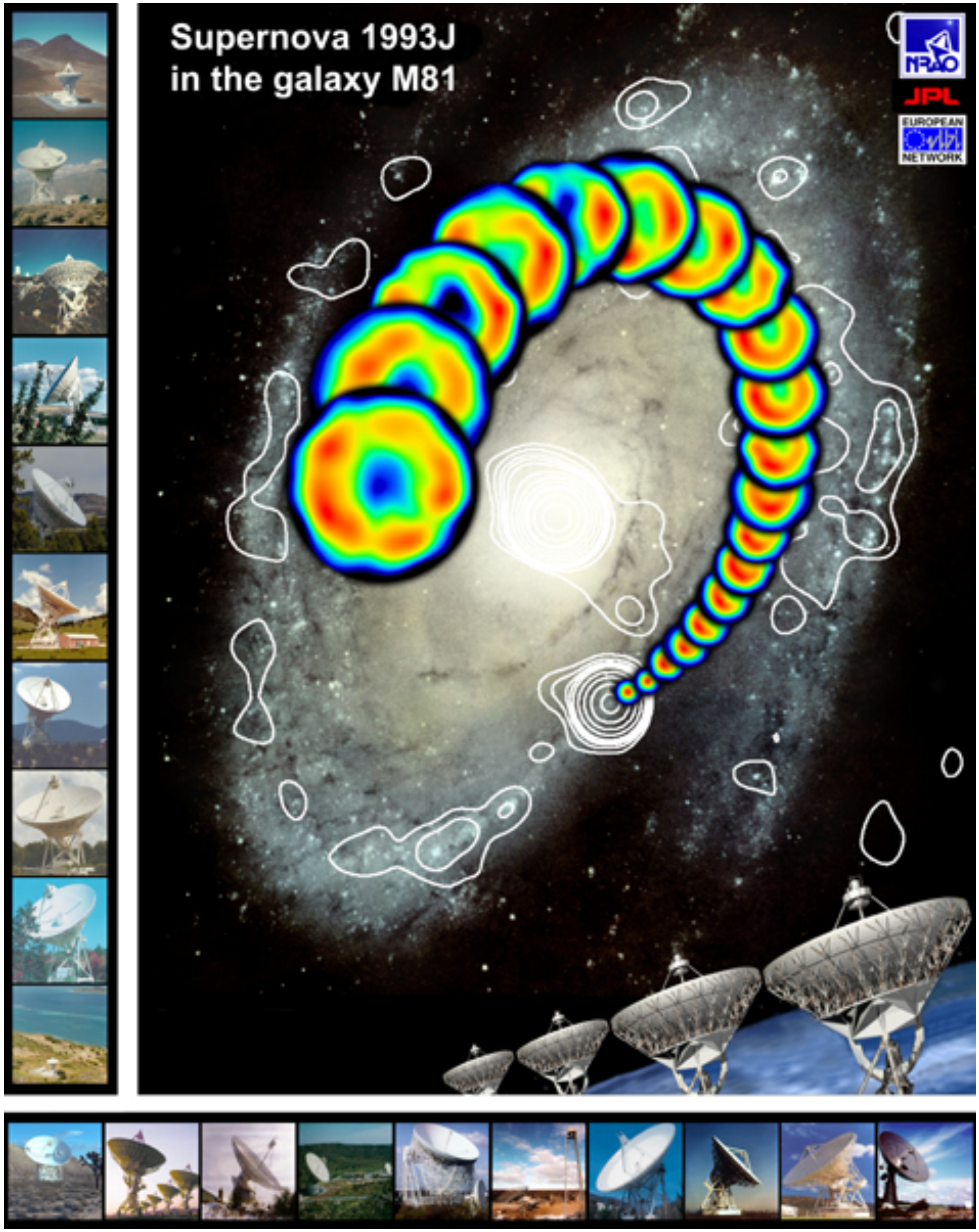
e Iron

Images (a) and (b) are computer simulations showing the chaotic flow of gas deep inside the star as it begins to explode as a supernova. This uneven flow helps account for the globs of iron and other heavy elements emitted from deep inside, as well as the lopsided distribution of all elements in the supernova remnant, as shown in (c), (d), and (e). These three pictures are X-ray images of supernova remnant Cassiopeia A taken by *Chandra* at different wavelengths.

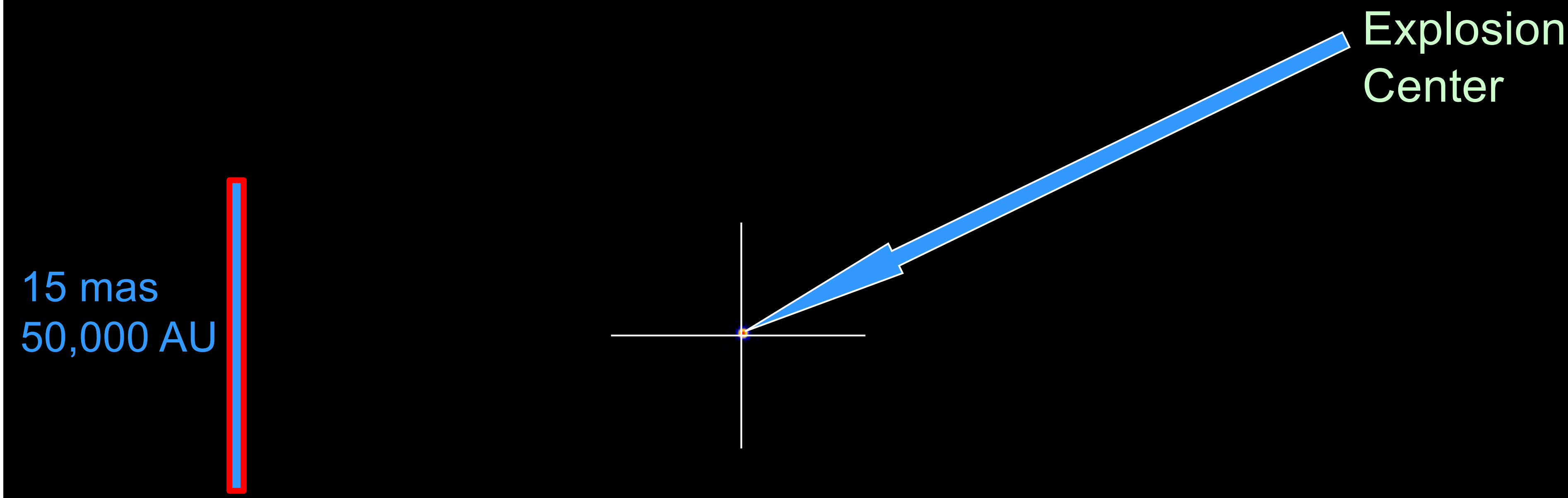


(a) Spiral galaxy M81 (b) Before the explosion (c) After the explosion

Figure 20-16
Universe, Eighth Edition


Supernova 1993J
in the galaxy M81
NRAO
JPL
EUROPEAN
NETWORK

Movie of SN1993J

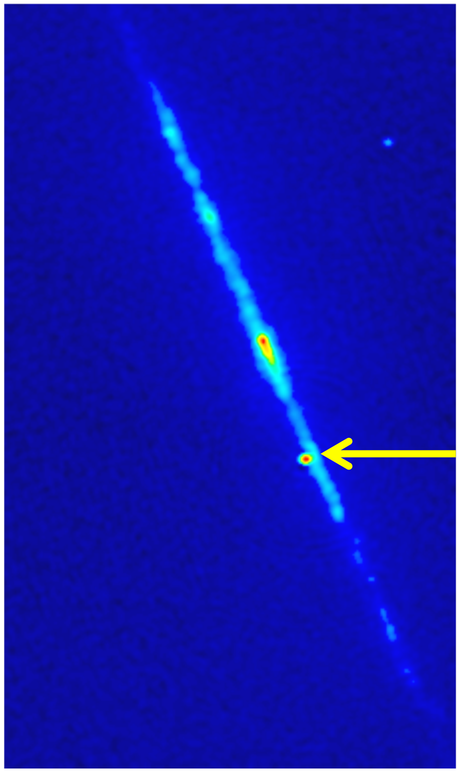


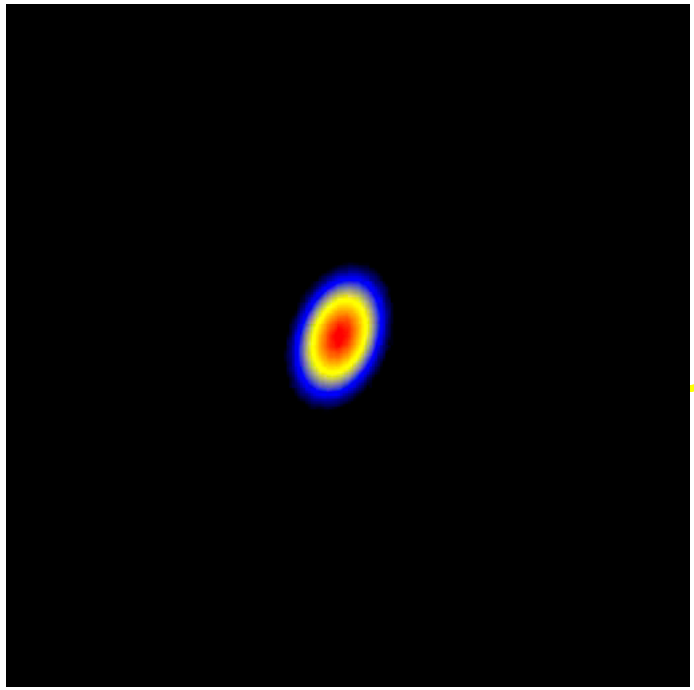
from t = 50d (r=200 AU) to t = 22 yr (r=40,000 AU)

Free download: www.yorku.ca/bartel

Evolution of SN1986J

NGC 891





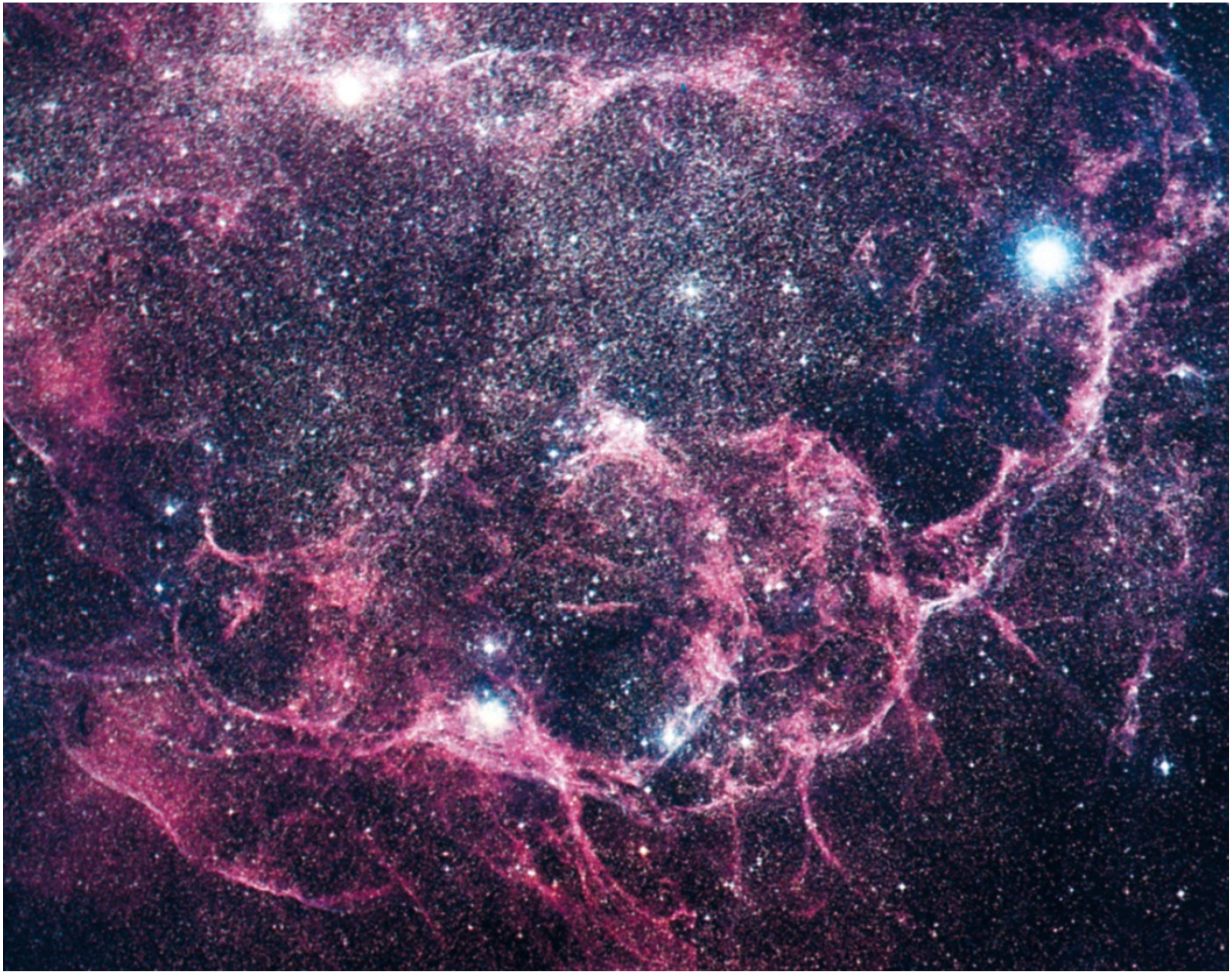
Youngest Neutron Star or Black Hole with 200 L_{Crab}?

Or dense shell condensation?

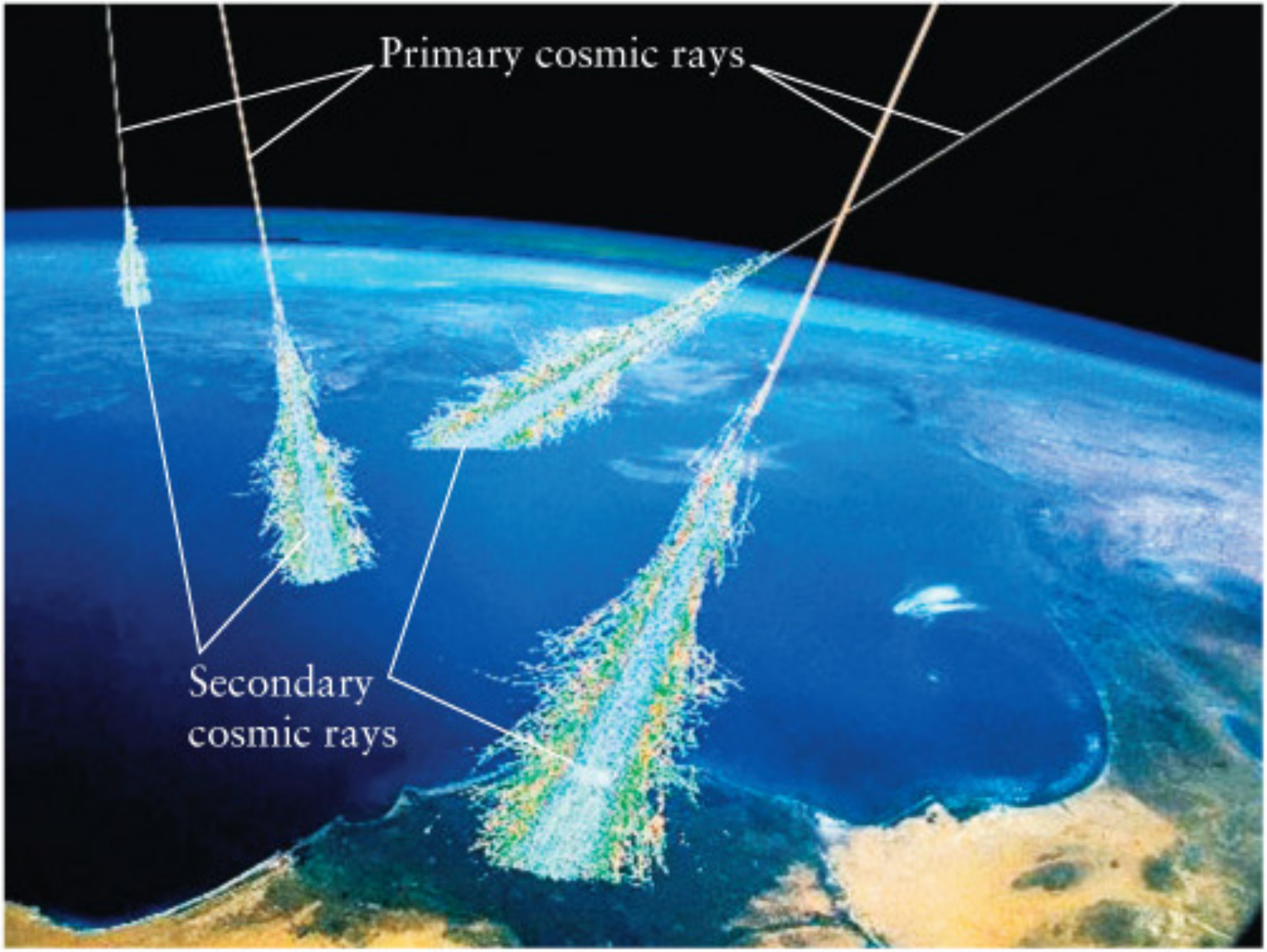
1 mas

from *t* = 3 yr to *t* = 25 yr

The Gum Nebula



Cosmic Ray Shower




Neutron Stars

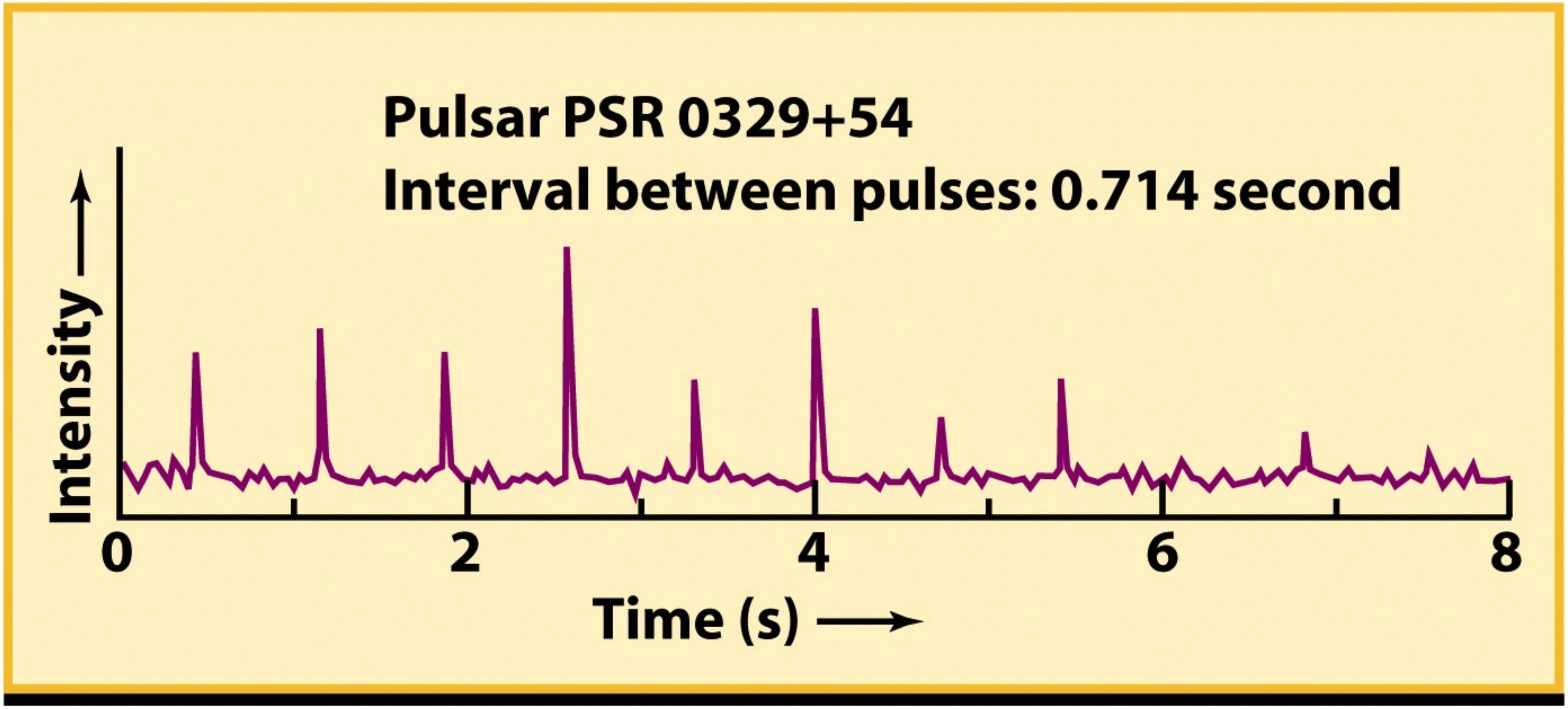


Figure 21-2
Universe, Eighth Edition

Why a pulsar must be a neutron star

- Radius of neutron star: 10 km
- Circumference of neutron star: ~60 km
- Fastest rotation from pulse period: 700/s
- Surface rotation speed: 42,000 km/s
- Escape velocity from WD: 1,500 km/s
- Escape velocity from NS: 150,000 km/s

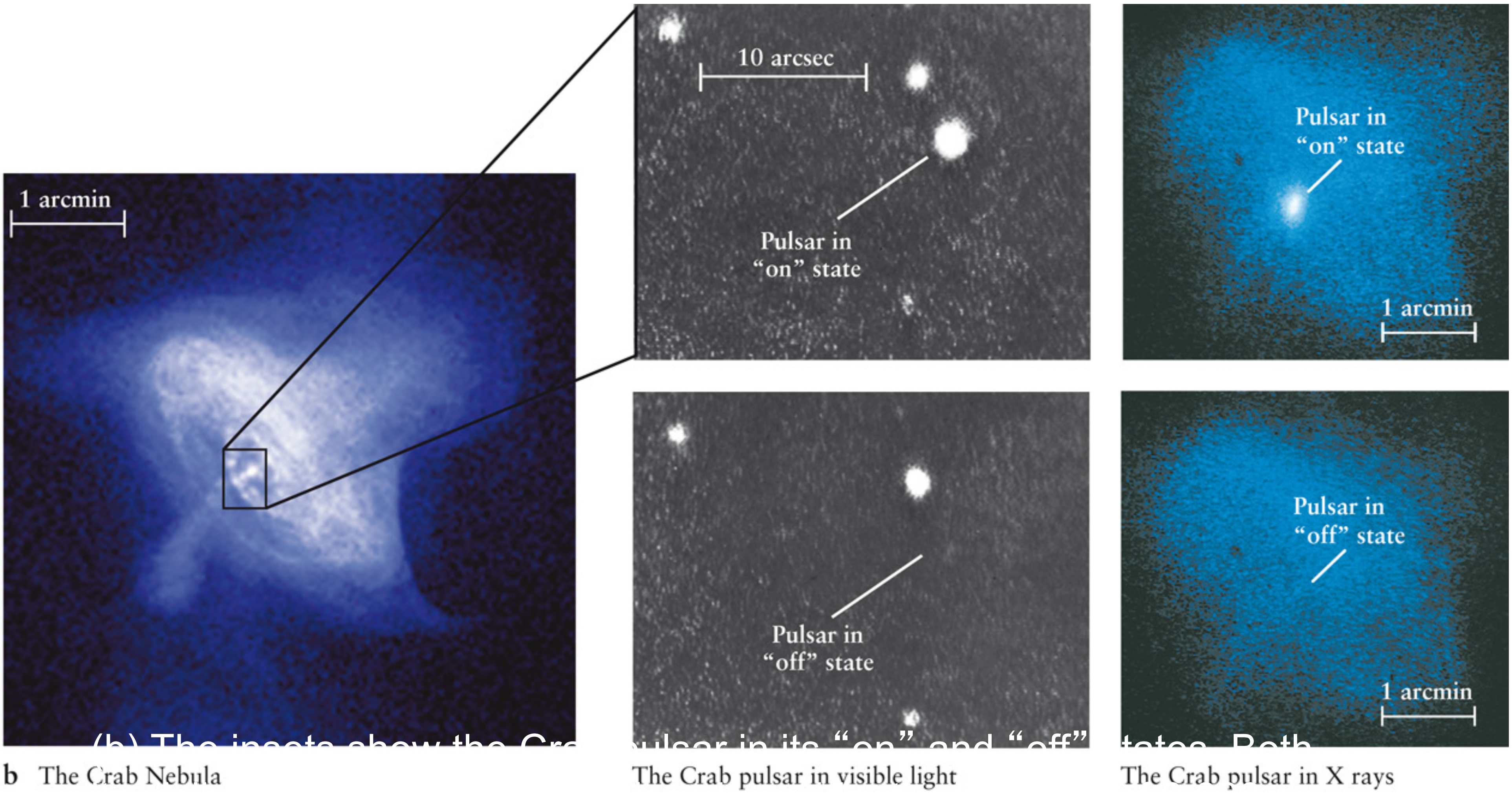
→ Everything less compact than a NS would disintegrate

Escape velocity

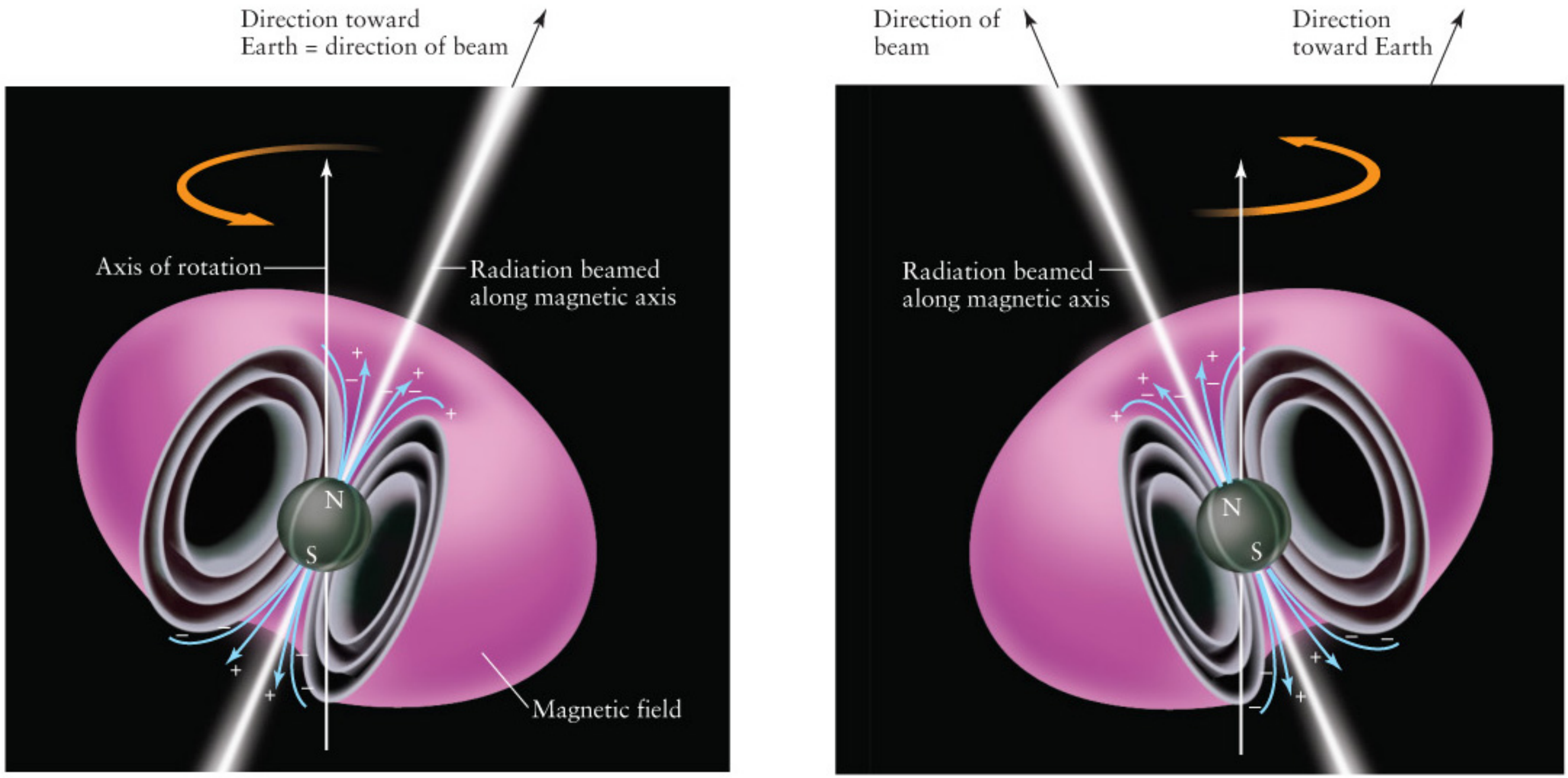
- The velocity that must be acquired by a body to just escape, i.e., to have zero total energy, is called the *escape velocity*. By setting $E_k + E_p = 0$, we find:

$$v^2_{escape} = 2\ G\ m\ /\ r$$

The Crab Nebula and Pulsar



A Rotating, Magnetized Neutron Star



a One of the beams from the rotating neutron star is aimed toward Earth: We detect a pulse of radiation.

b Half a rotation later, neither beam is aimed toward Earth: We detect that the radiation is "off."

The relativity theories

- Special relativity (1905)
- General relativity (1916)

Einstein in school (~ 1892)

Bern with clock tower

Franz Anneler, Malerweg 9 (Längg.) u. Vereinsdruckerei.
Guss von Wälzchen für Mimiographen.

Vermischtes

Privatstunden in
Mathematik u. Physik
für Studierende und Schüler erteilt gründlichst
Albert Einstein, Inhaber des eidgen. polyt. Fachlehrerdiploms,
Gerechtigkeitsgasse 32, 1. Stock.
Probestunden gratis. 4977°

...wo es von einer tüchti- ...allen häuslichen Ar- ...würde. 4978 ...erstraße 7a, 2. St.

Mädchen, ...ig kochen kann, sucht ...reau. 5049

...son empfiehlt sich als ...eiderin, ...arbeit zu Hause od. ...sse im Bureau. 5051

..., junger ...häftigung. ...2. Glock...

...es, deutsch... Stelle al... möglich i... ...äre Behandlung gro-

...mit zwei Schaufenste... seit vier Jahren ein Geschäft en gros mit gutem Erfolge b... ist wegen Uebernahm... Geschäftes zu über... Gefl. schriftliche ... Nr. 5028 an das B...

...kaufmann sucht, um sein Geschäft besser zu entwickeln,

Vorzeich... malen und sticken von

Einstein's advertisement in the Berne paper, 1902: "Private lessons in mathematics and physics. . . ."

Free trial lesson

Einstein at the Swiss Patent Office

1905 - Miraculous Year

``A storm broke loose in my mind’’

- Einstein wrote three fundamental papers.

The 1st paper claimed that light must sometimes behave like a stream of particles with discrete energies, ”quanta.”

The 2nd paper offered an experimental test for the theory of heat. Atoms do exist!

The 3rd paper addressed a central puzzle for physicists of the day – the connection between electromagnetic theory and ordinary motion – and solved it using the "principle of relativity."

Special Relativity

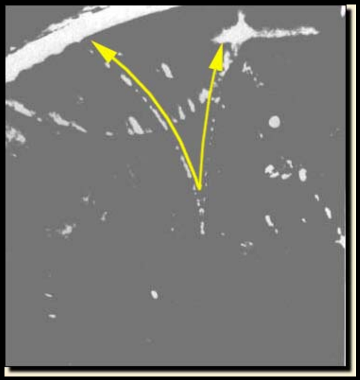
- 1905
- The laws of physics are the same for all inertial observers
- c=constant

→ Clocks slow down when in motion

→ Objects contract when in motion

→ $E=mc^2$

→ Spacetime



Einstein's list of conditions to his wife Mileva

before they separated in 1914

Mileva with Eduard and Hans Albert 1914

"You make sure ... that I receive my three meals regularly in my room.
You are neither to expect intimacy nor to reproach me in any way."

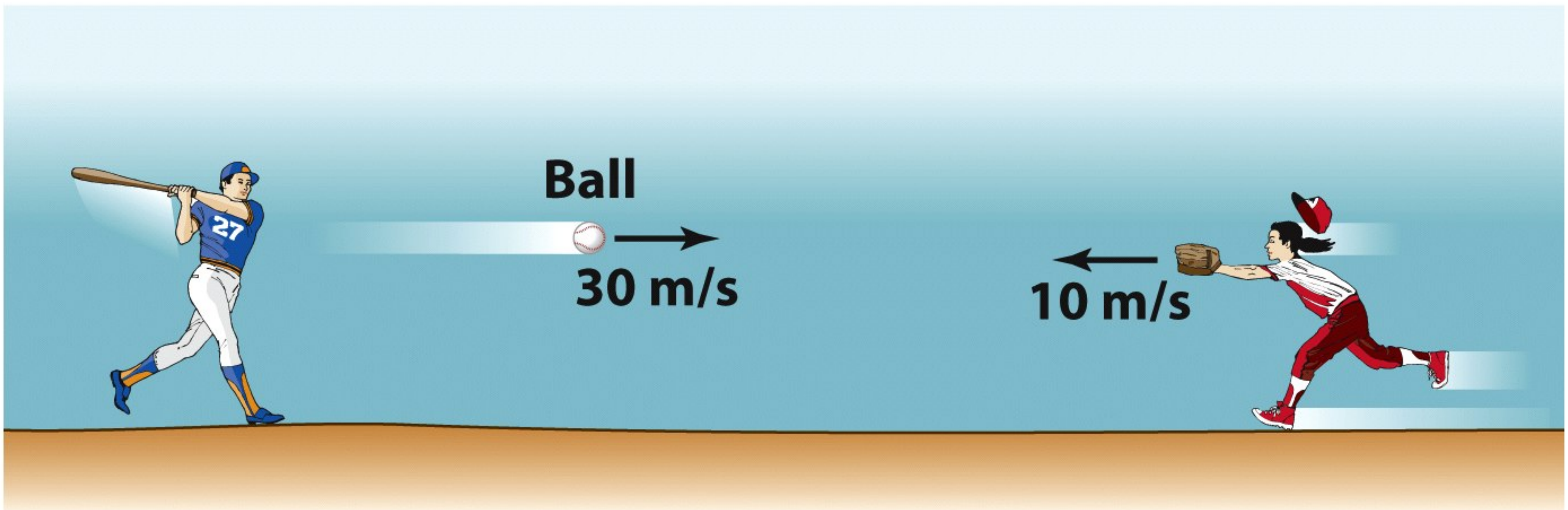


Figure 22-1a
Universe, Eighth Edition

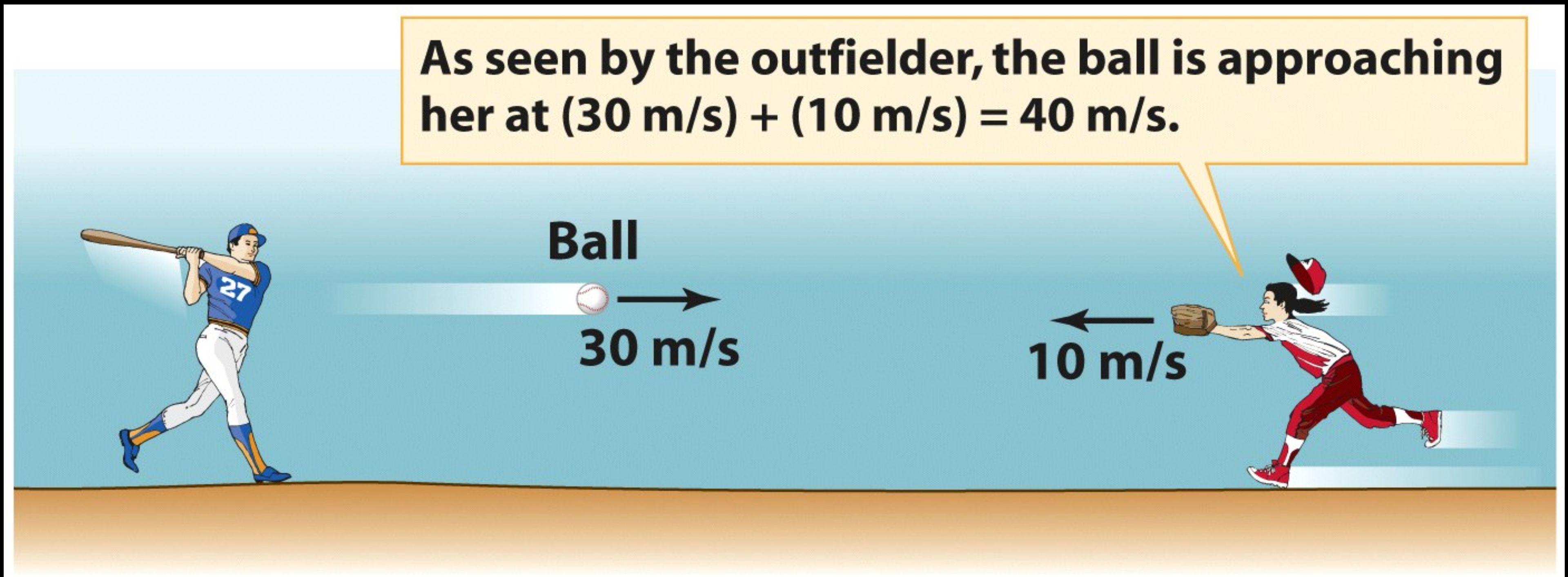


Figure 22-1a
Universe, Eighth Edition

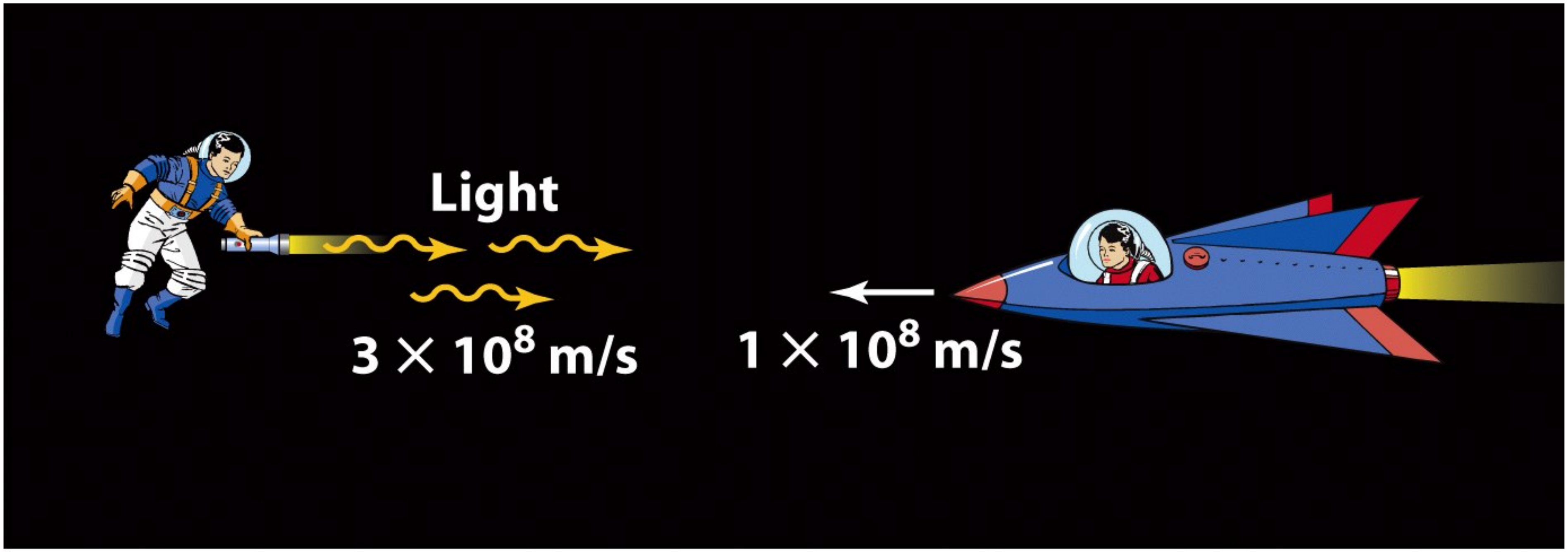


Figure 22-1b
Universe, Eighth Edition

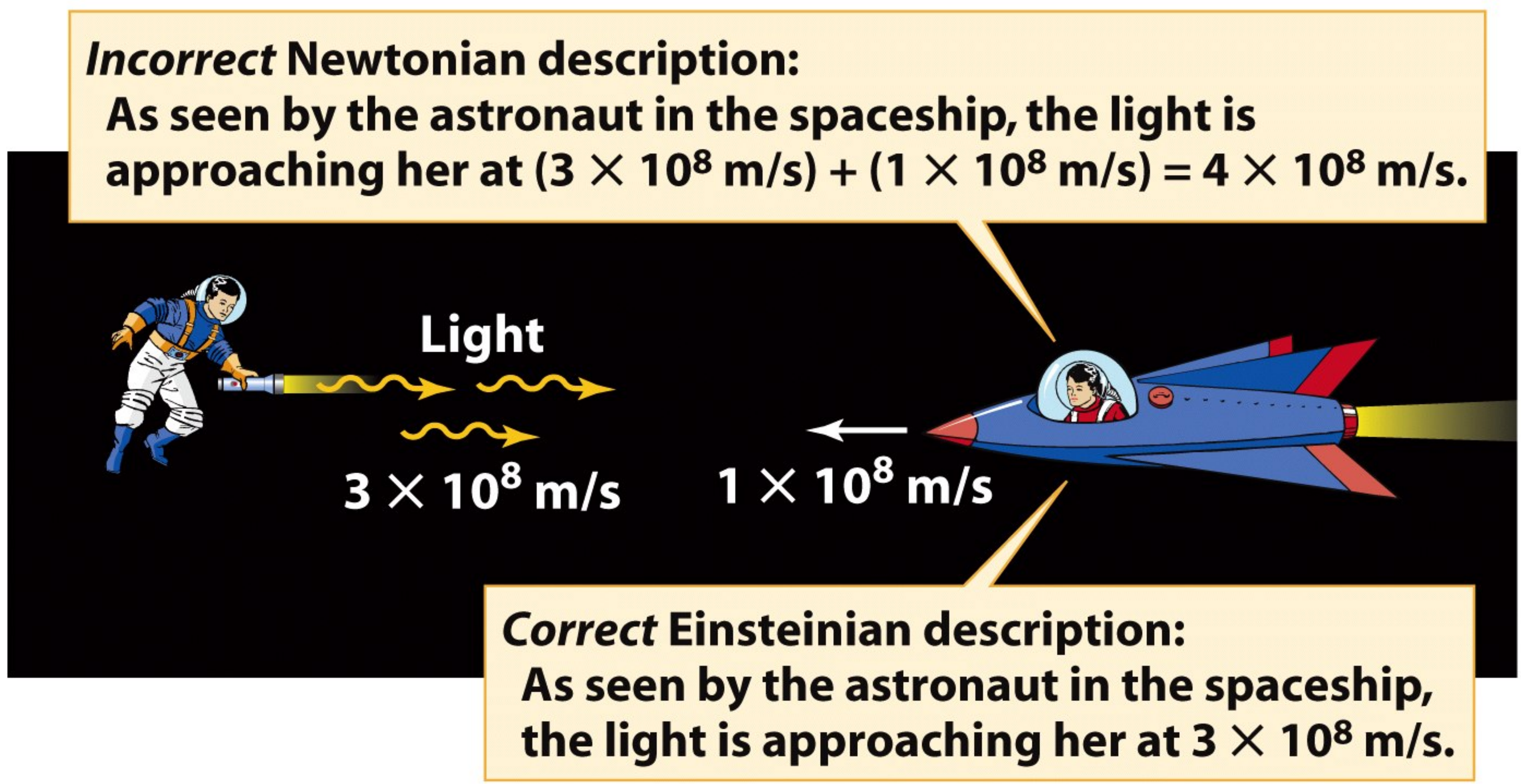
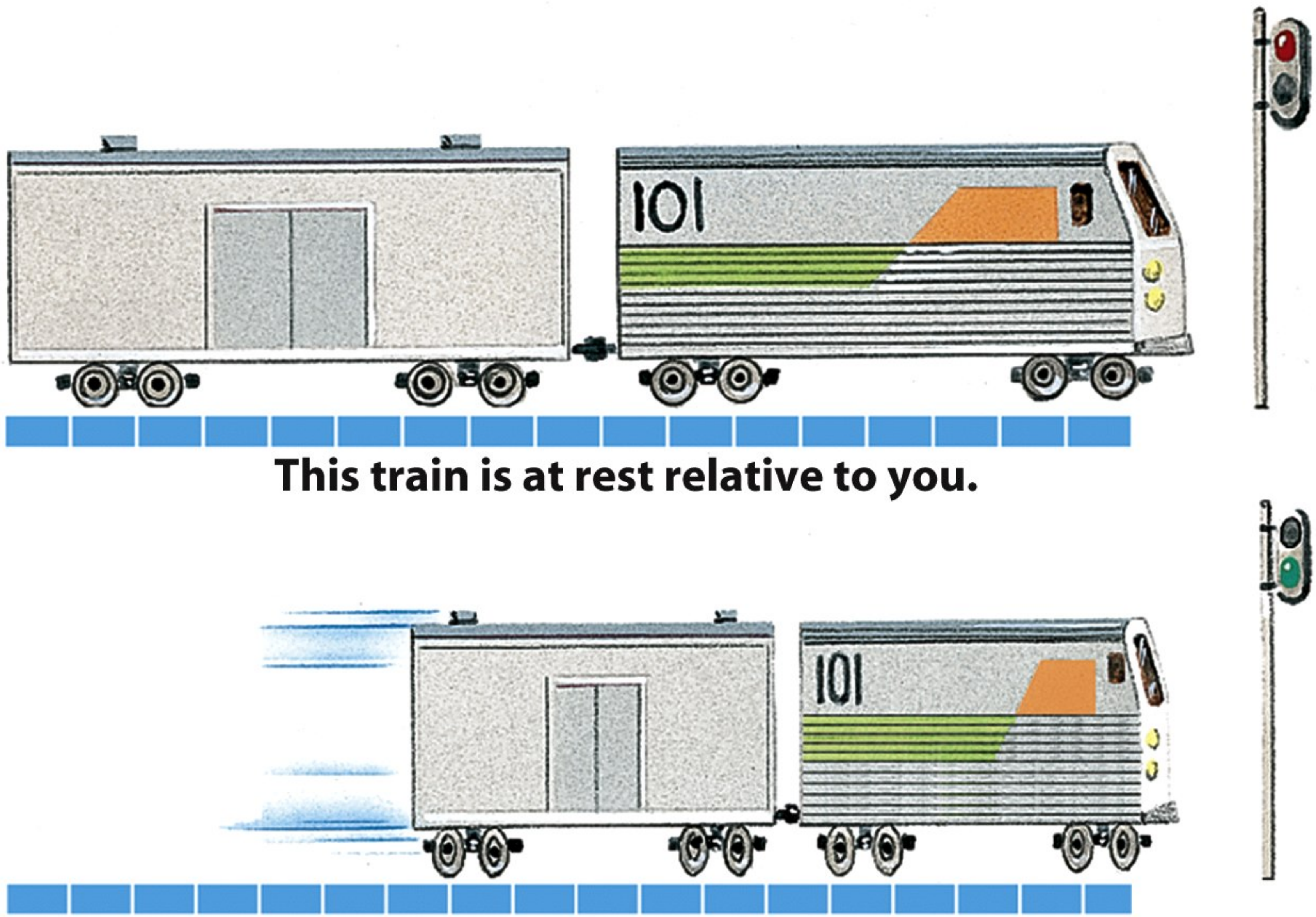


Figure 22-1b
Universe, Eighth Edition



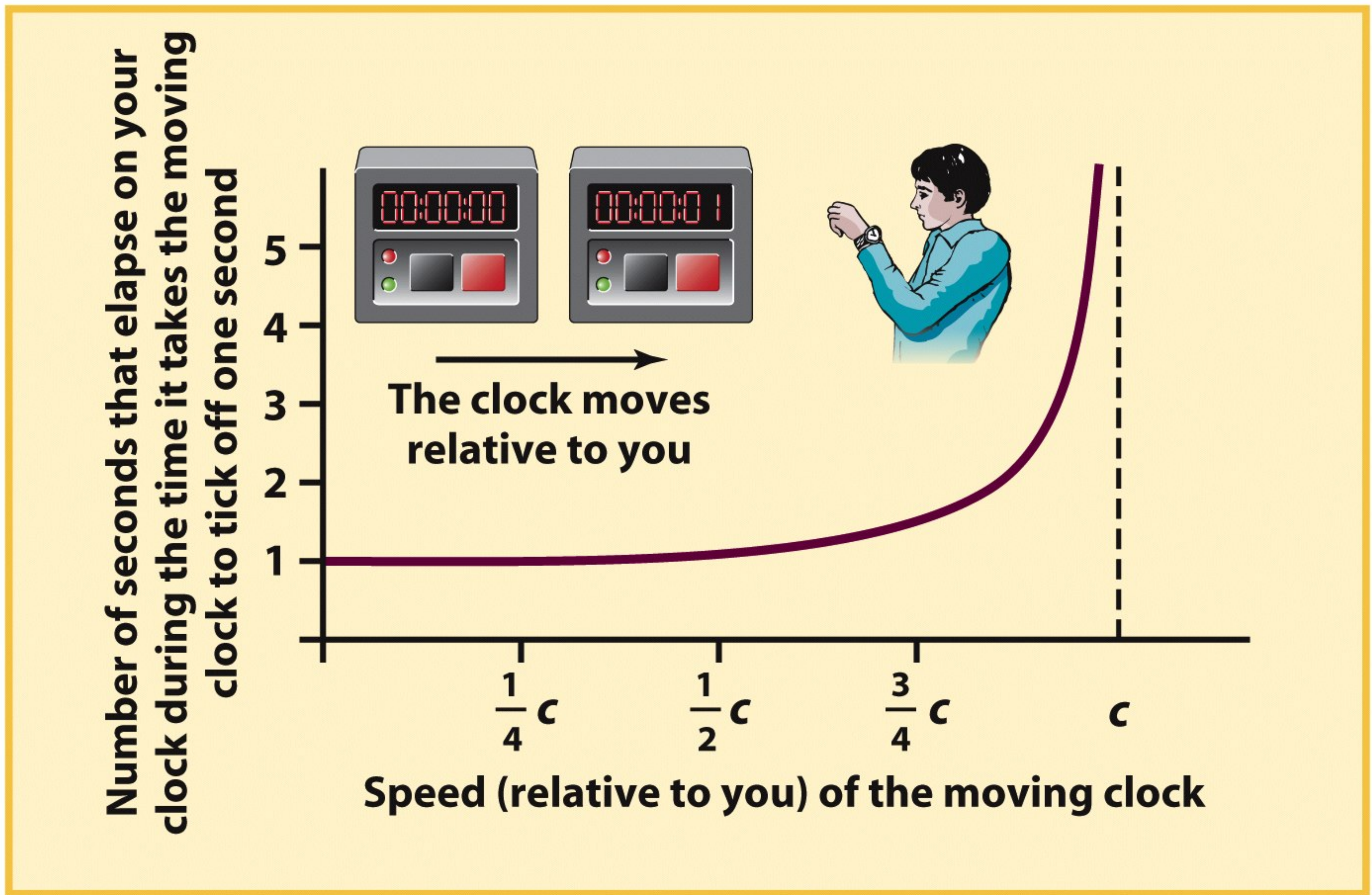
This train is at rest relative to you.

The same train is now moving relative to you.

Length contraction

$$L=L_0[1-(v/c)^2]^{-1/2}$$

Figure 22-2a
Universe, Eighth Edition

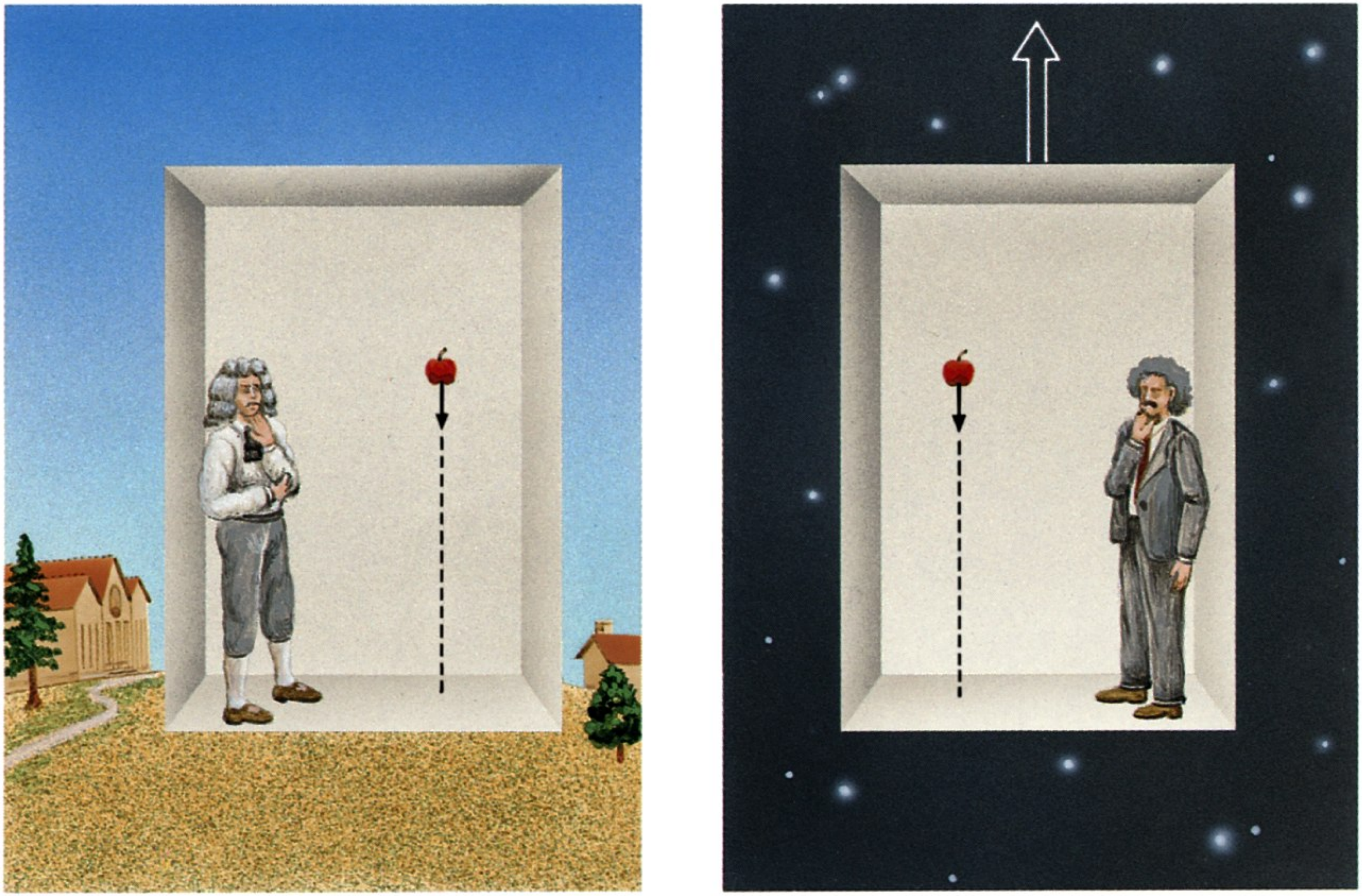


Time dilation

$$T=T_0[1-(v/c)^2]^{-1/2}$$

Figure 22-2b

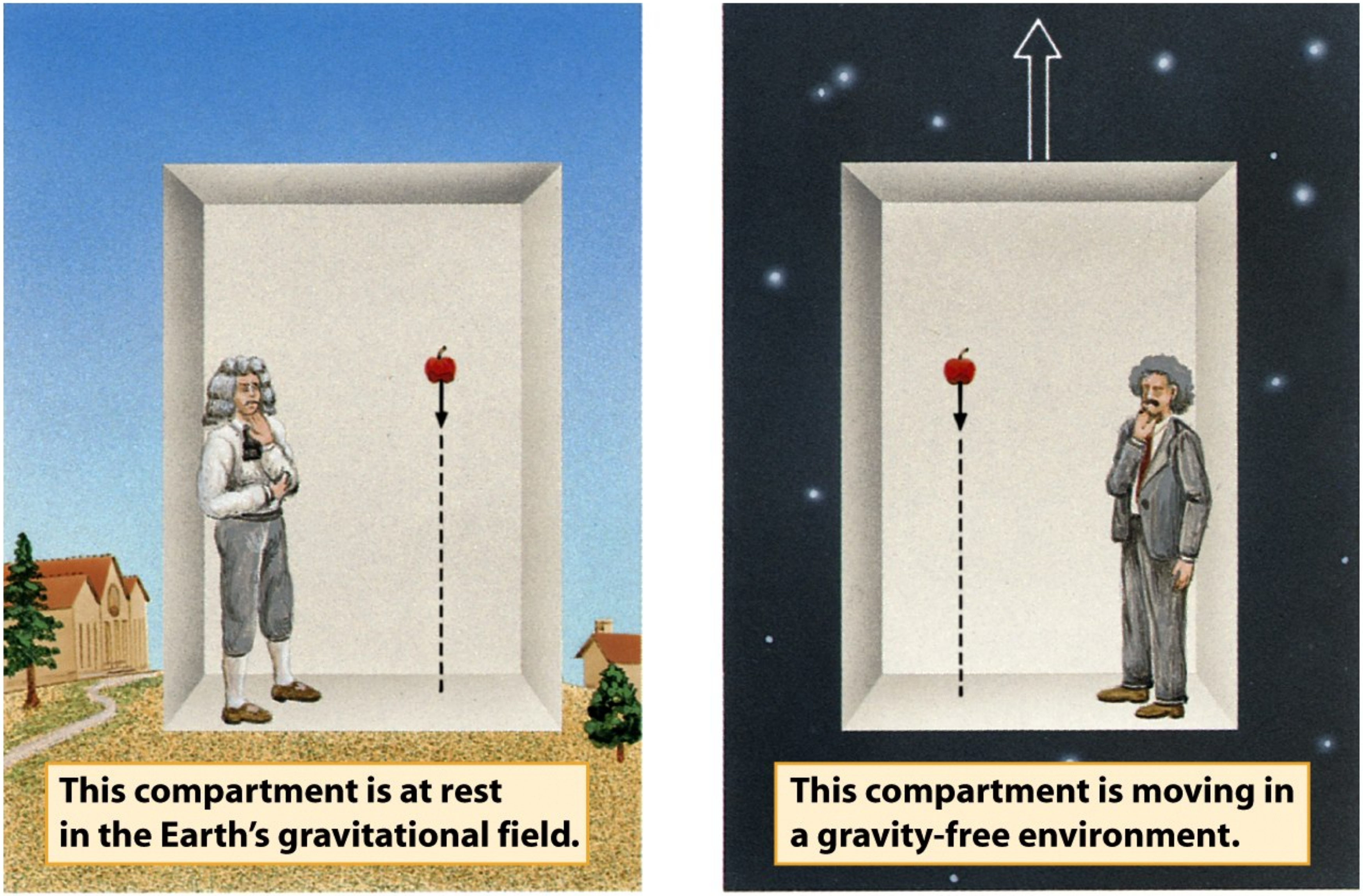
Universe, Eighth Edition



(a) The apple hits the floor of the compartment because the Earth's gravity accelerates the apple downward.

(b) The apple hits the floor of the compartment because the compartment accelerates upward.

Figure 22-3
Universe, Eighth Edition



(a) The apple hits the floor of the compartment because the Earth's gravity accelerates the apple downward.

(b) The apple hits the floor of the compartment because the compartment accelerates upward.

Figure 22-3
Universe, Eighth Edition

General Relativity

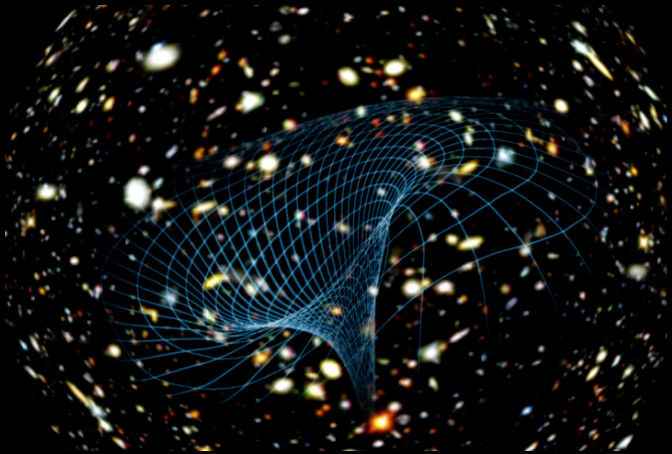
- 1916
- Special relativity + gravitation

→ Curved and twisted spacetime

$$G_{\mu\nu} - \Lambda g_{\mu\nu} = -8\pi G T_{\mu\nu}$$

→ Matter and energy tell spacetime how to curve and spacetime tells matter how to move

→ Big Bang, black holes, the Universe



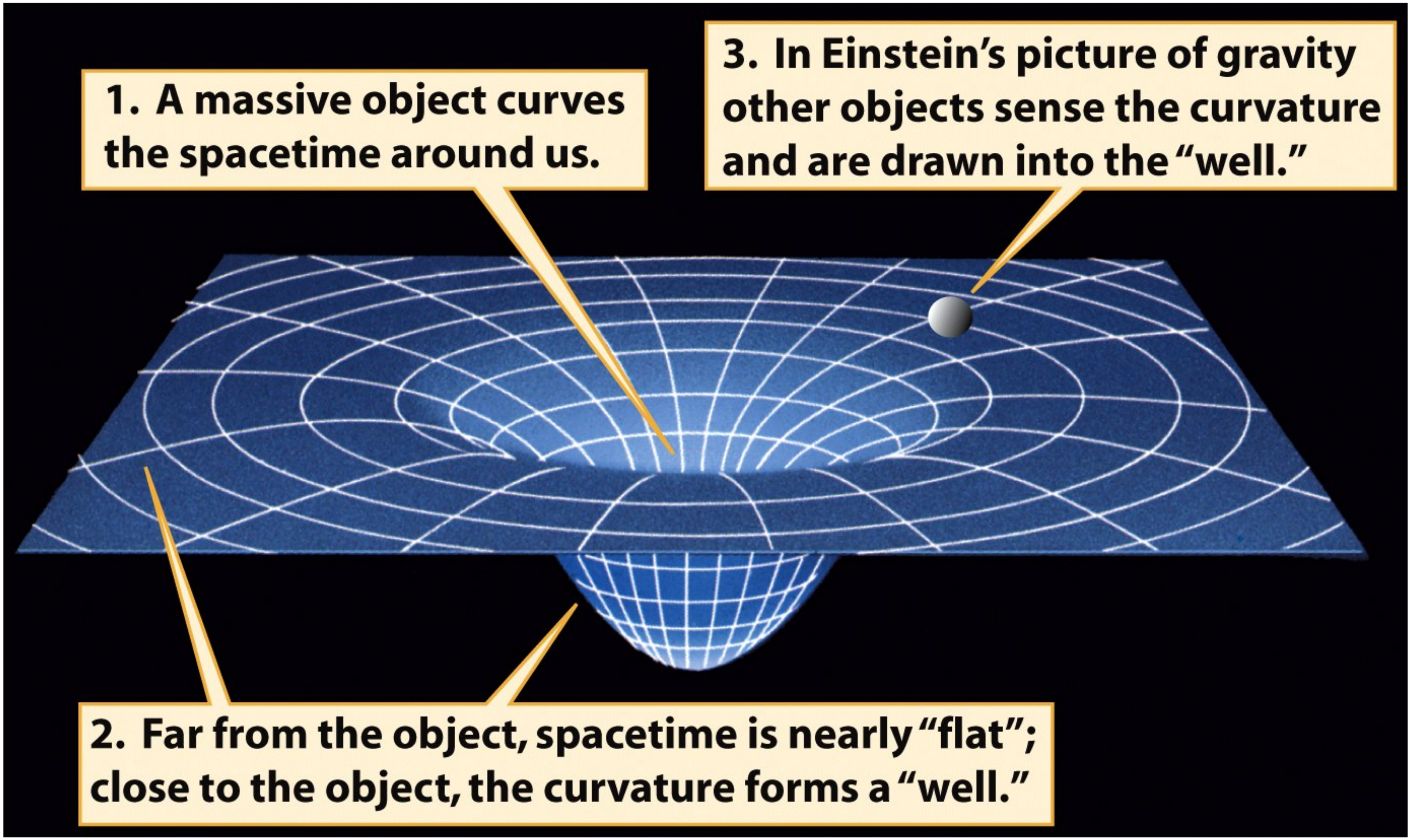
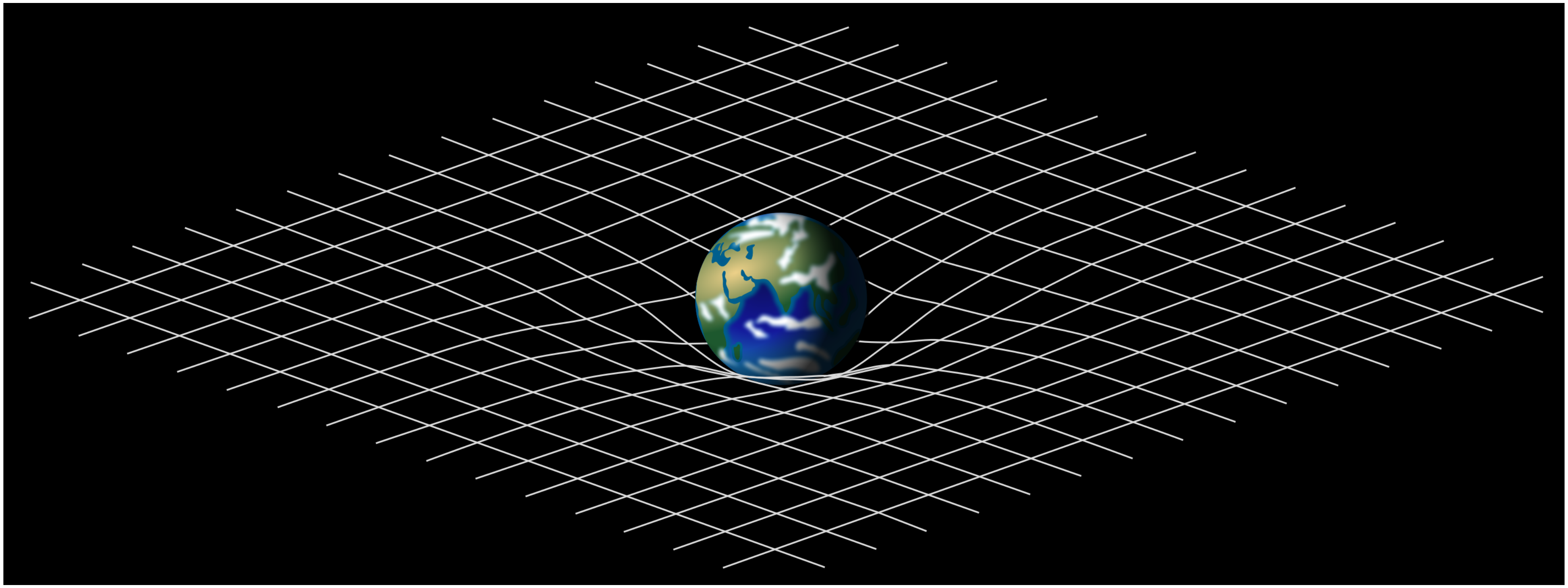


Figure 22-4
Universe, Eighth Edition

Gravity distorts spacetime



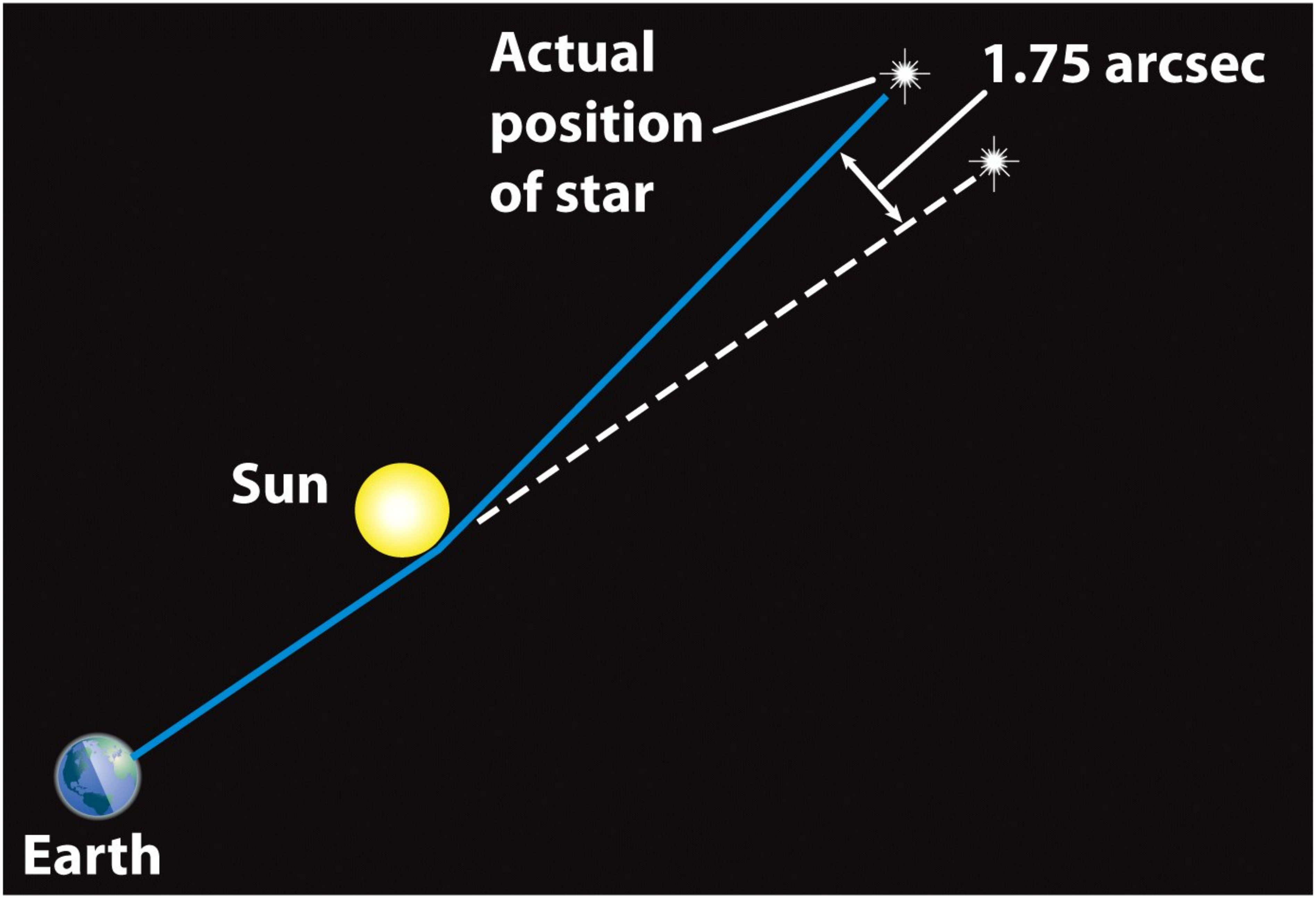
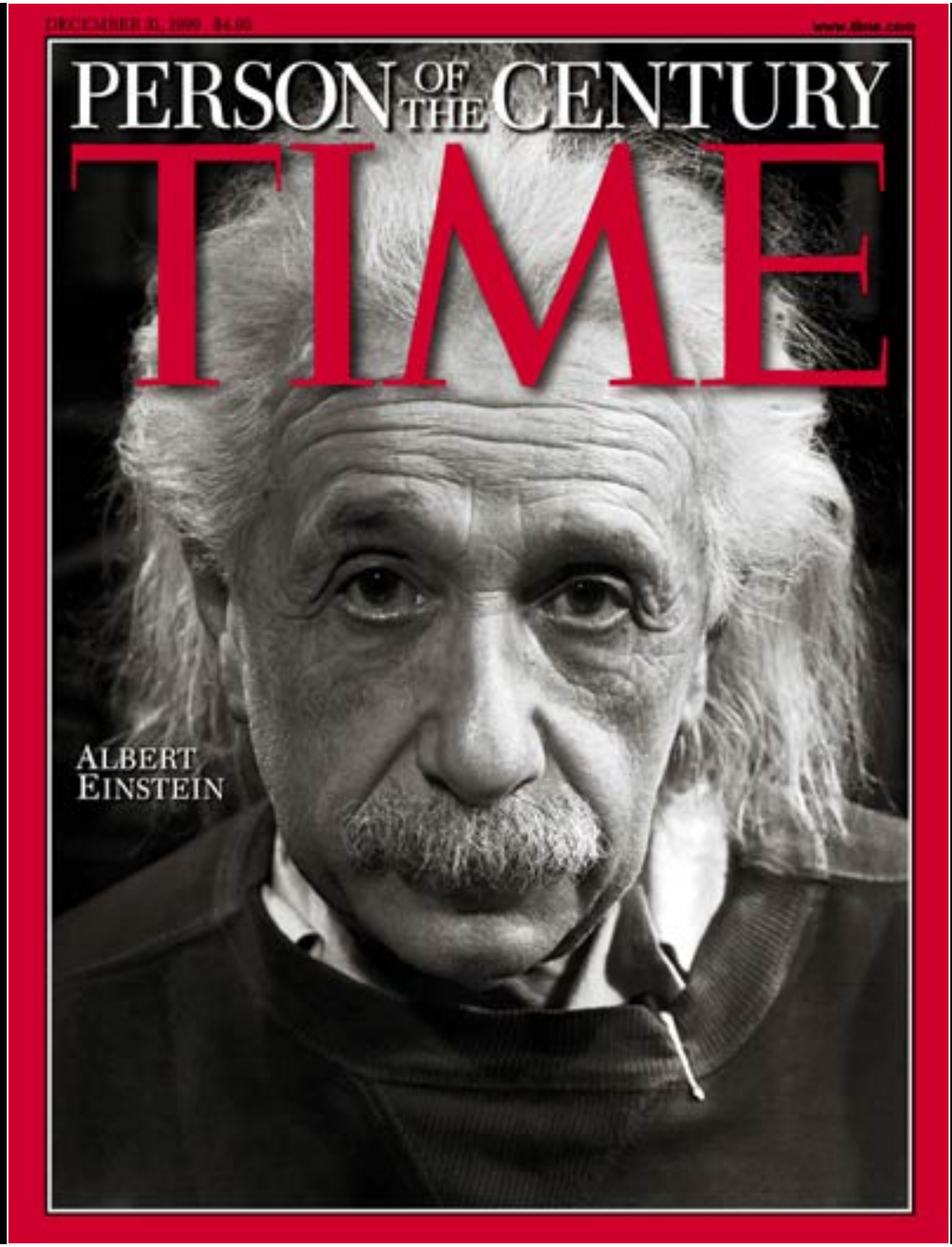


Figure 22-5
Universe, Eighth Edition

Einstein's first visit to the United States, in 1921

"I feel like a prima donna"


PERSON OF THE CENTURY
TIME
ALBERT
EINSTEIN

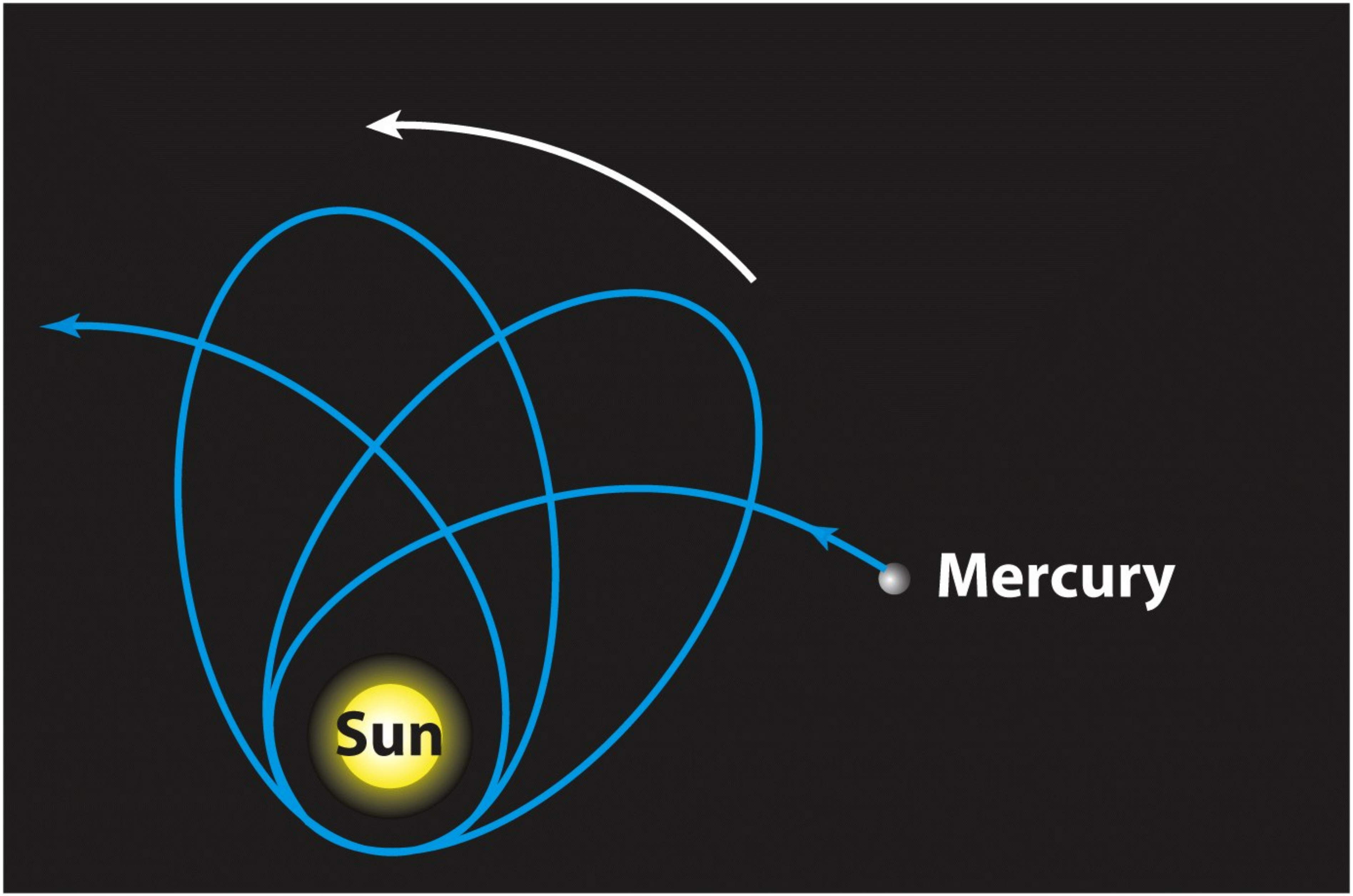
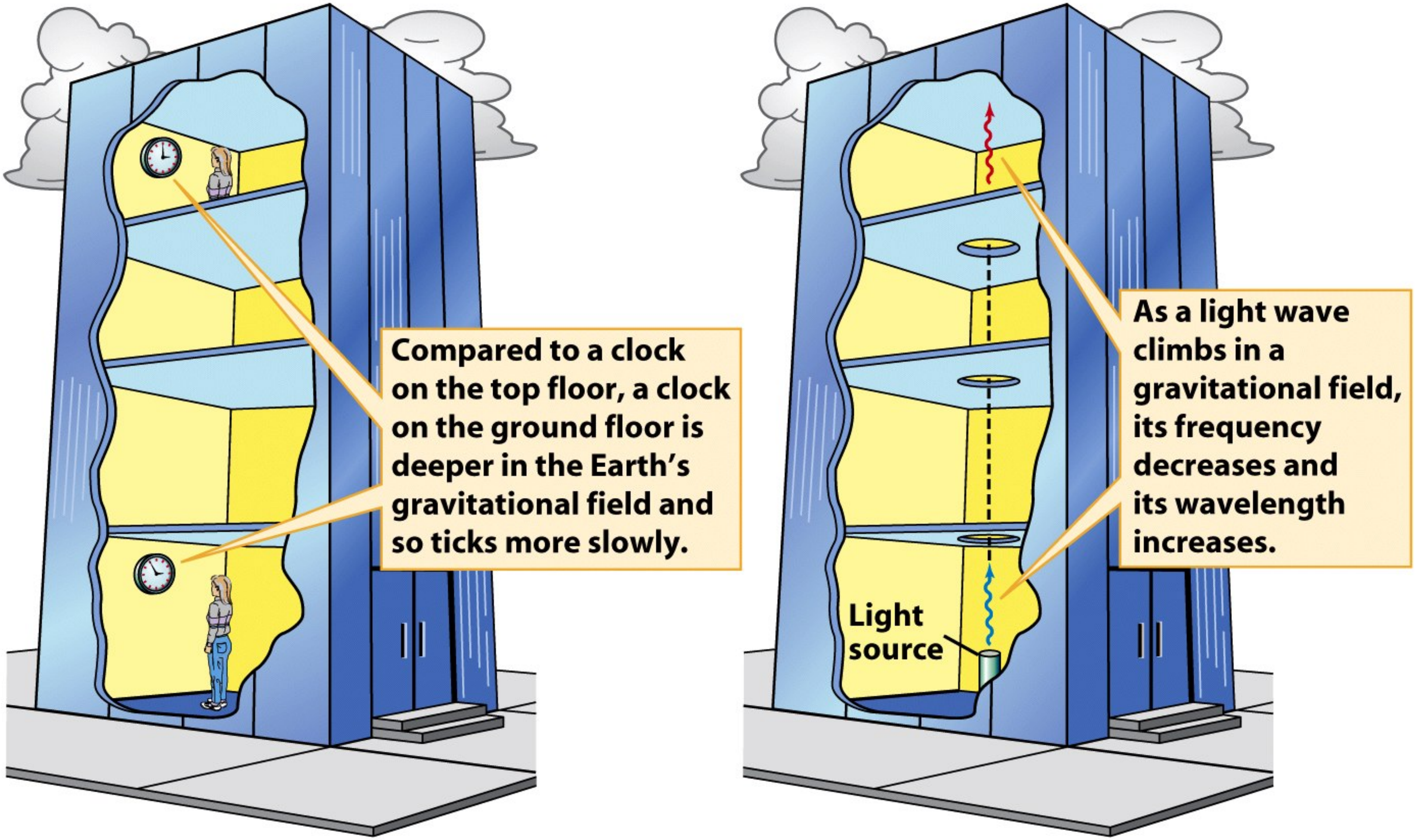


Figure 22-6
Universe, Eighth Edition



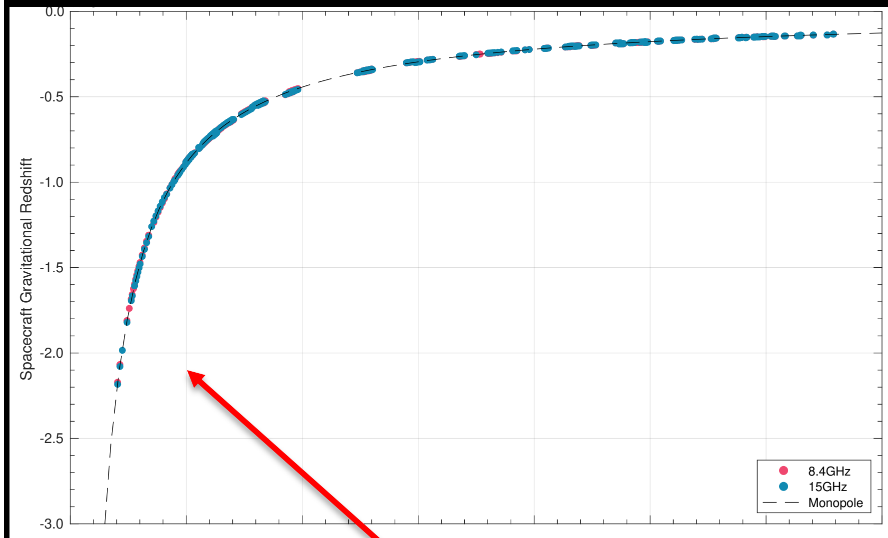
(a) The gravitational slowing of time

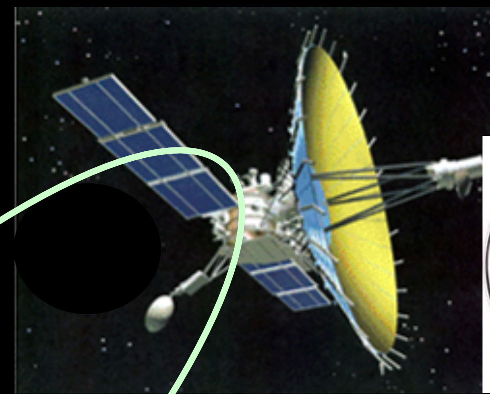
(b) The gravitational redshift

Figure 22-7
Universe, Eighth Edition

Probing gravitational redshift.
-Grad student Nelson Nunes

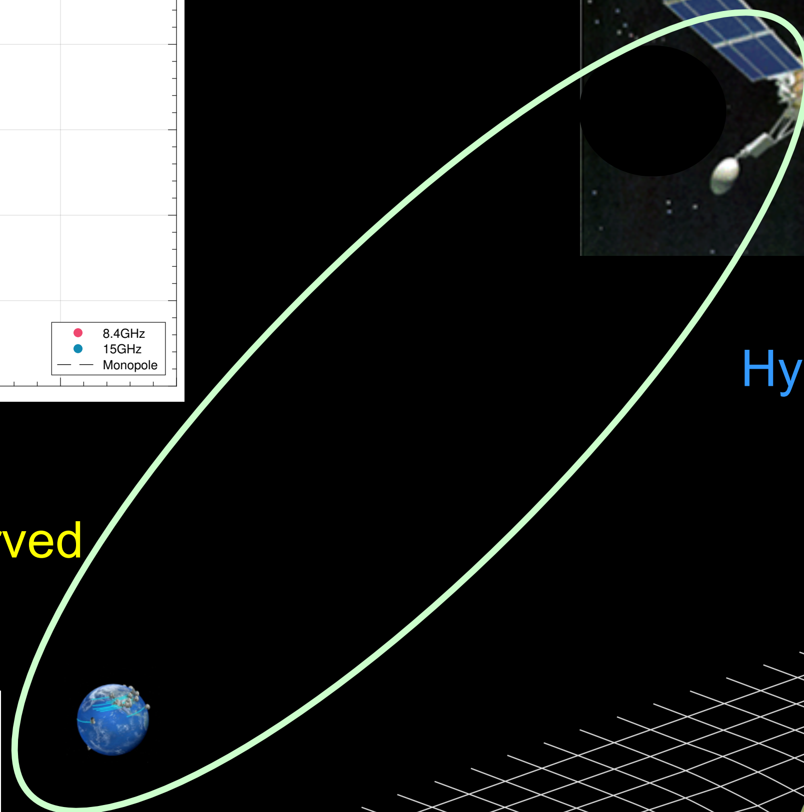
RadioAstron
Ground-space VLBI mission

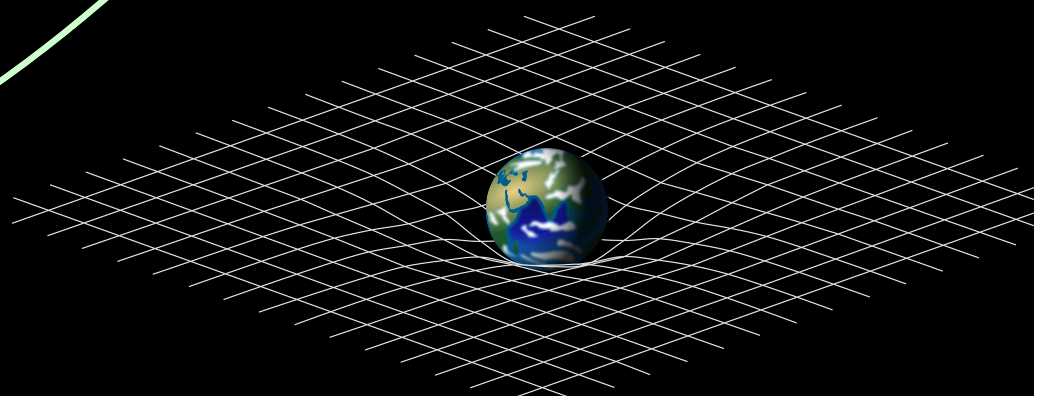




Hydrogen maser clock-2

Near Earth:
Time slows down, space is curved



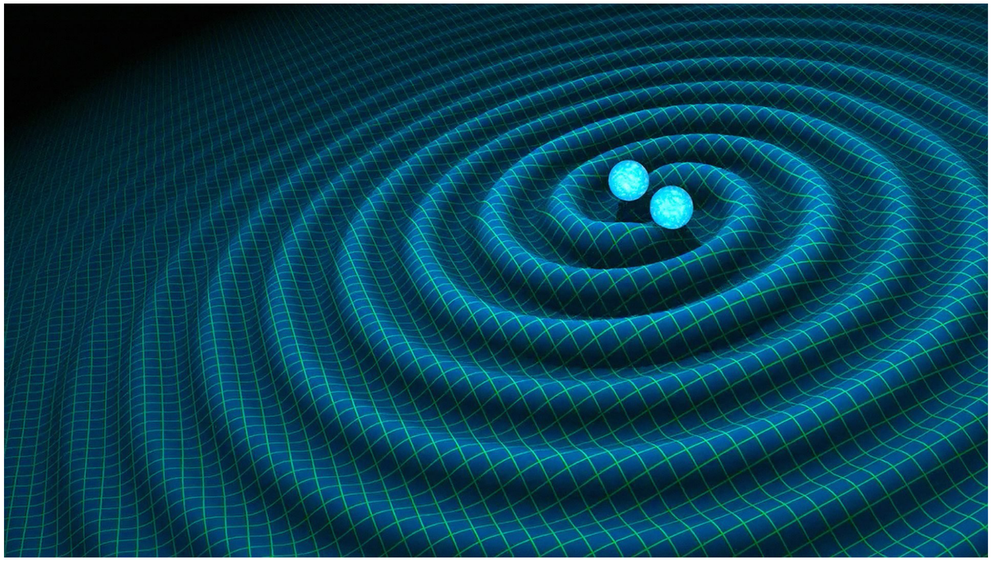


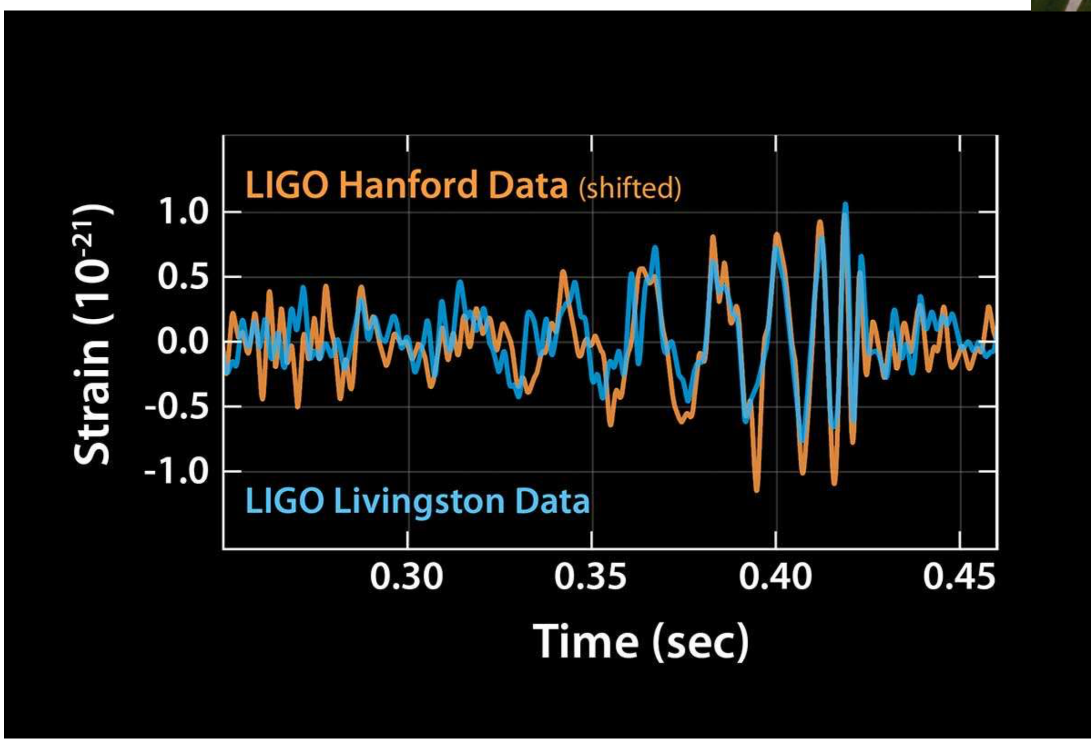
Hydrogen maser clock-1

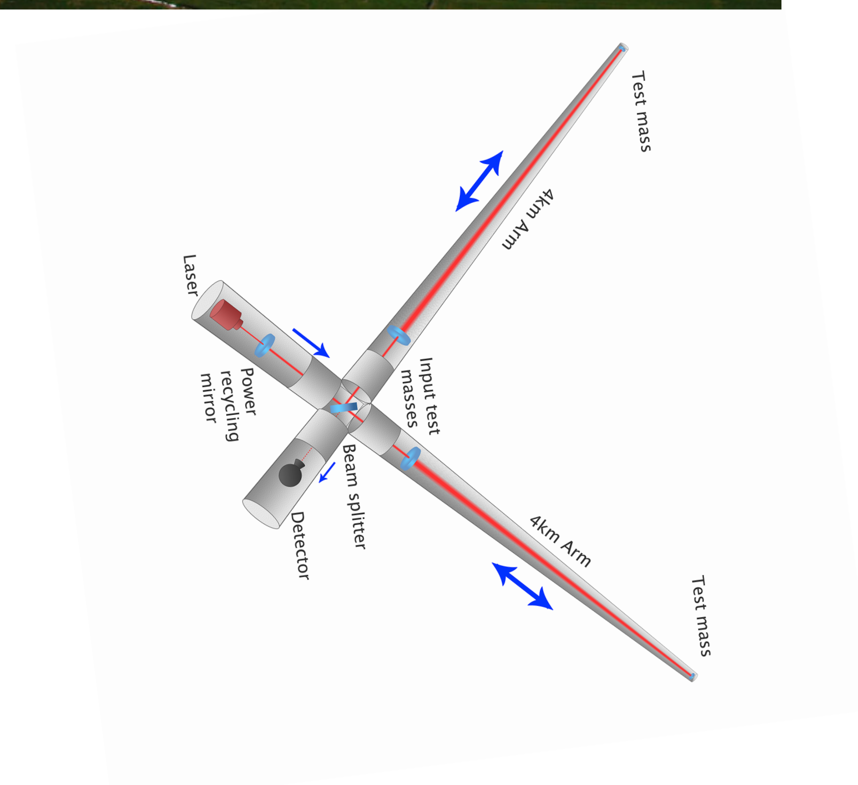
Wikipedia

Gravitational waves

two colliding black holes







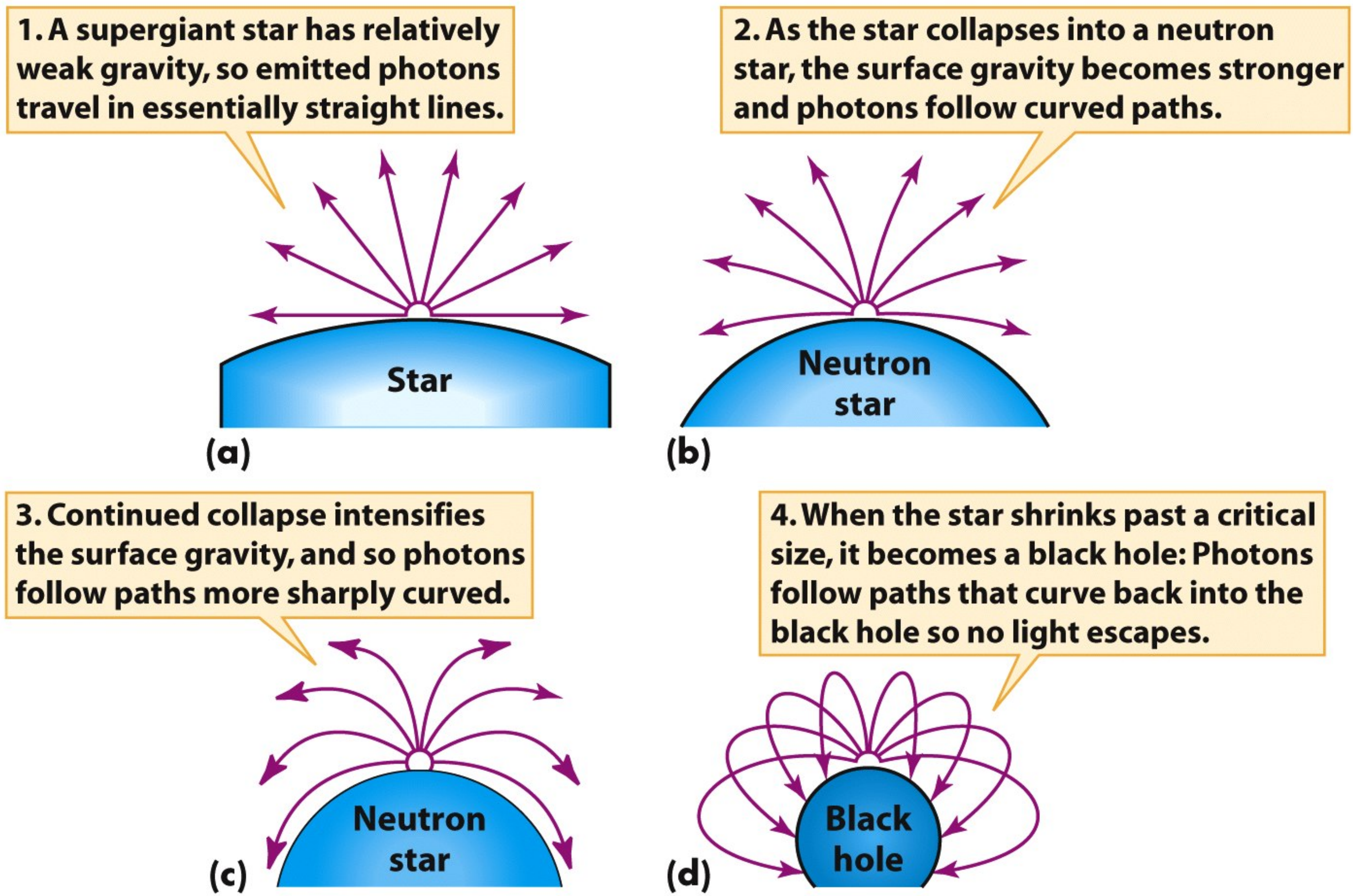
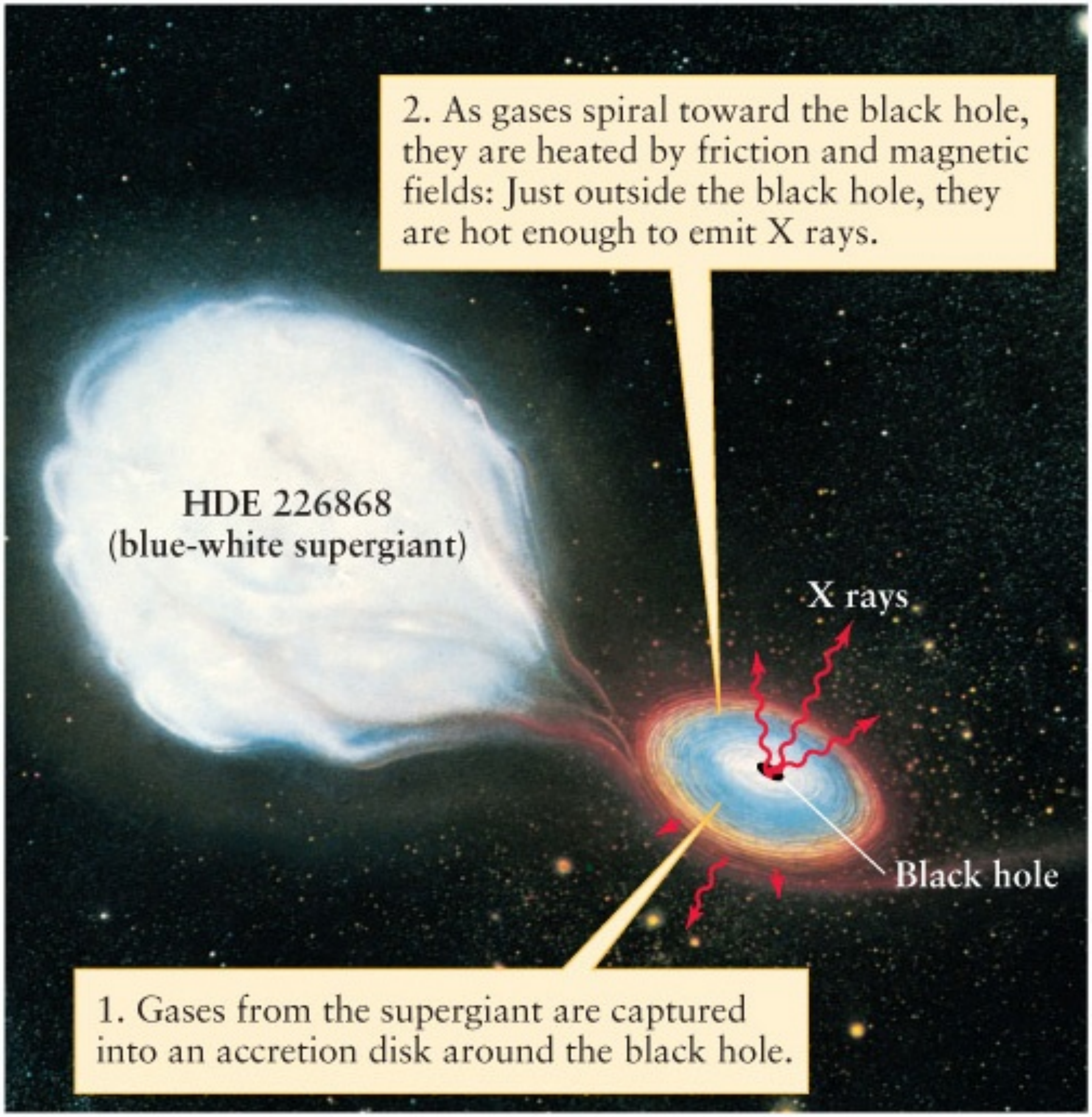
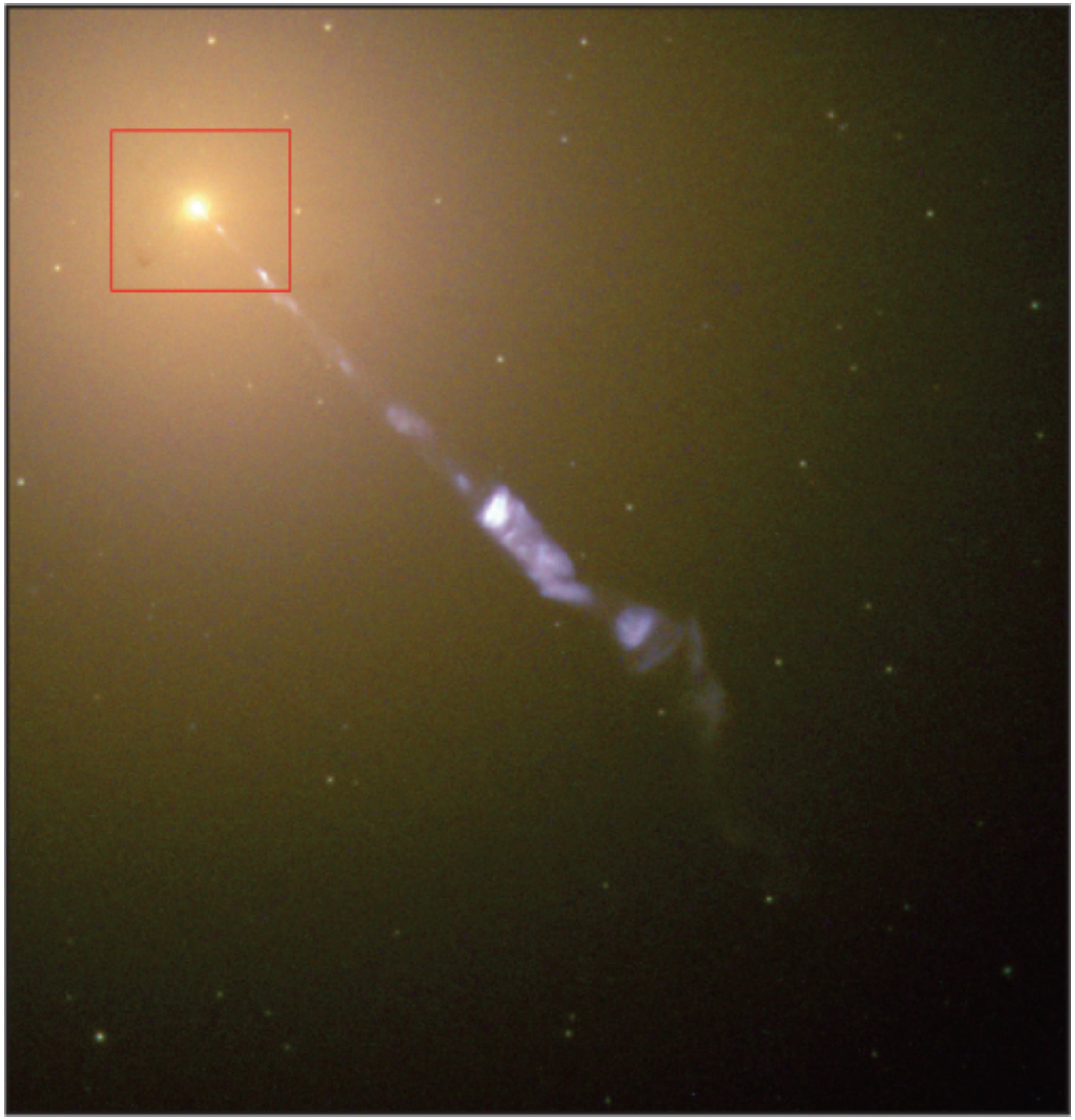


Figure 22-8
Universe, Eighth Edition

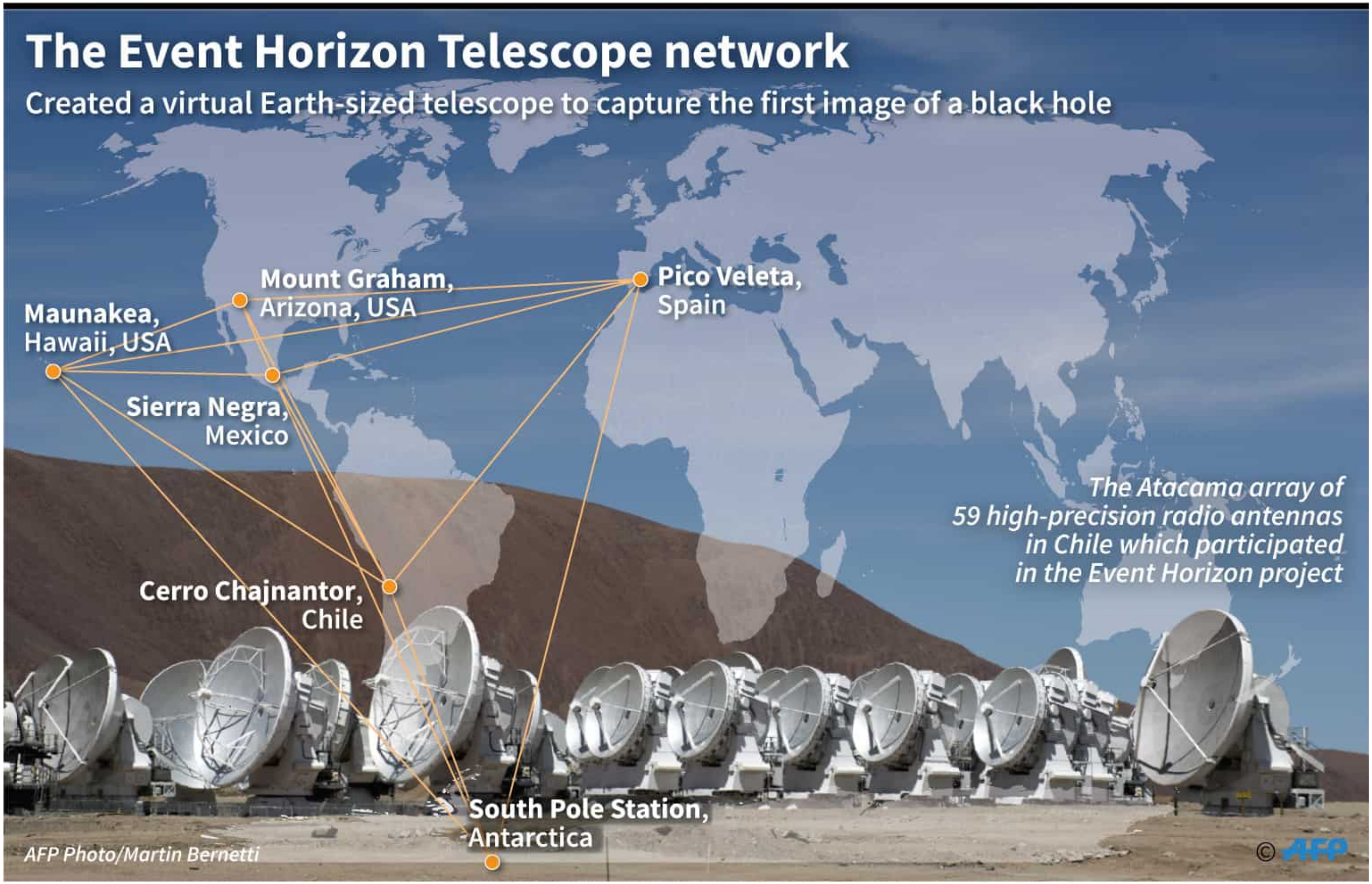
X Rays Generated by Accretion of Matter Near a Black Hole



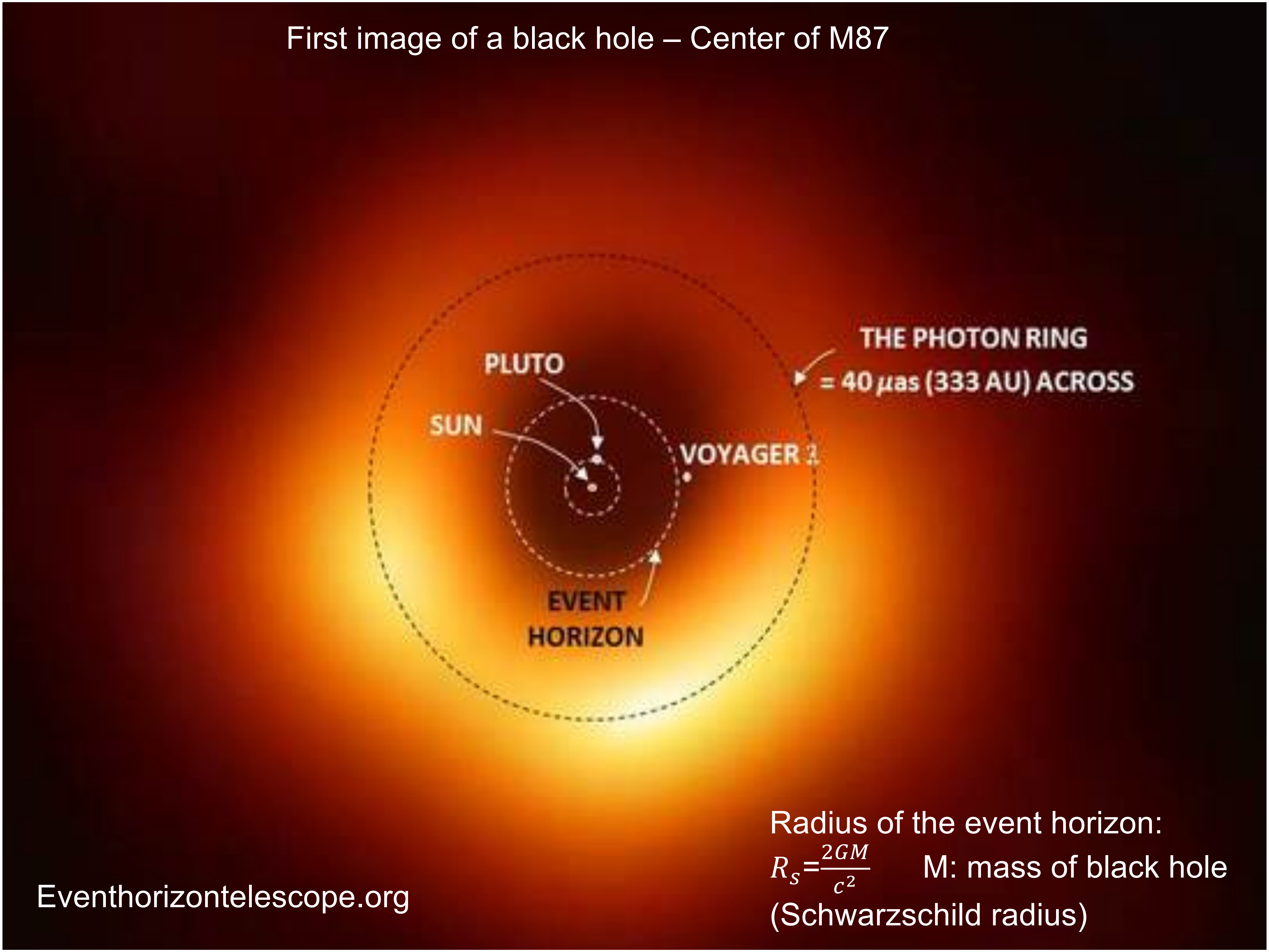
Supermassive Black Hole in the galaxy M87



The bright region in the center of galaxy M87 has stars and gas held in tight orbits by a black hole. M87's bright nucleus (center of the region in the white box) is only about the size of the solar system but it pulls on the nearby stars with so much force that astronomers calculate that it is a 6-billion-solar-mass black hole. One of the bright jets of gas shooting out perpendicular to the black hole's accretion disk is also visible.


The Event Horizon Telescope network
Created a virtual Earth-sized telescope to capture the first image of a black hole
Mount Graham, Arizona, USA
Pico Veleta, Spain
Maunakea, Hawaii, USA
Sierra Negra, Mexico
Cerro Chajnantor, Chile
The Atacama array of 59 high-precision radio antennas in Chile which participated in the Event Horizon project
South Pole Station, Antarctica
AFP Photo/Martin Bernetti
© AFP

First image of a black hole – Center of M87

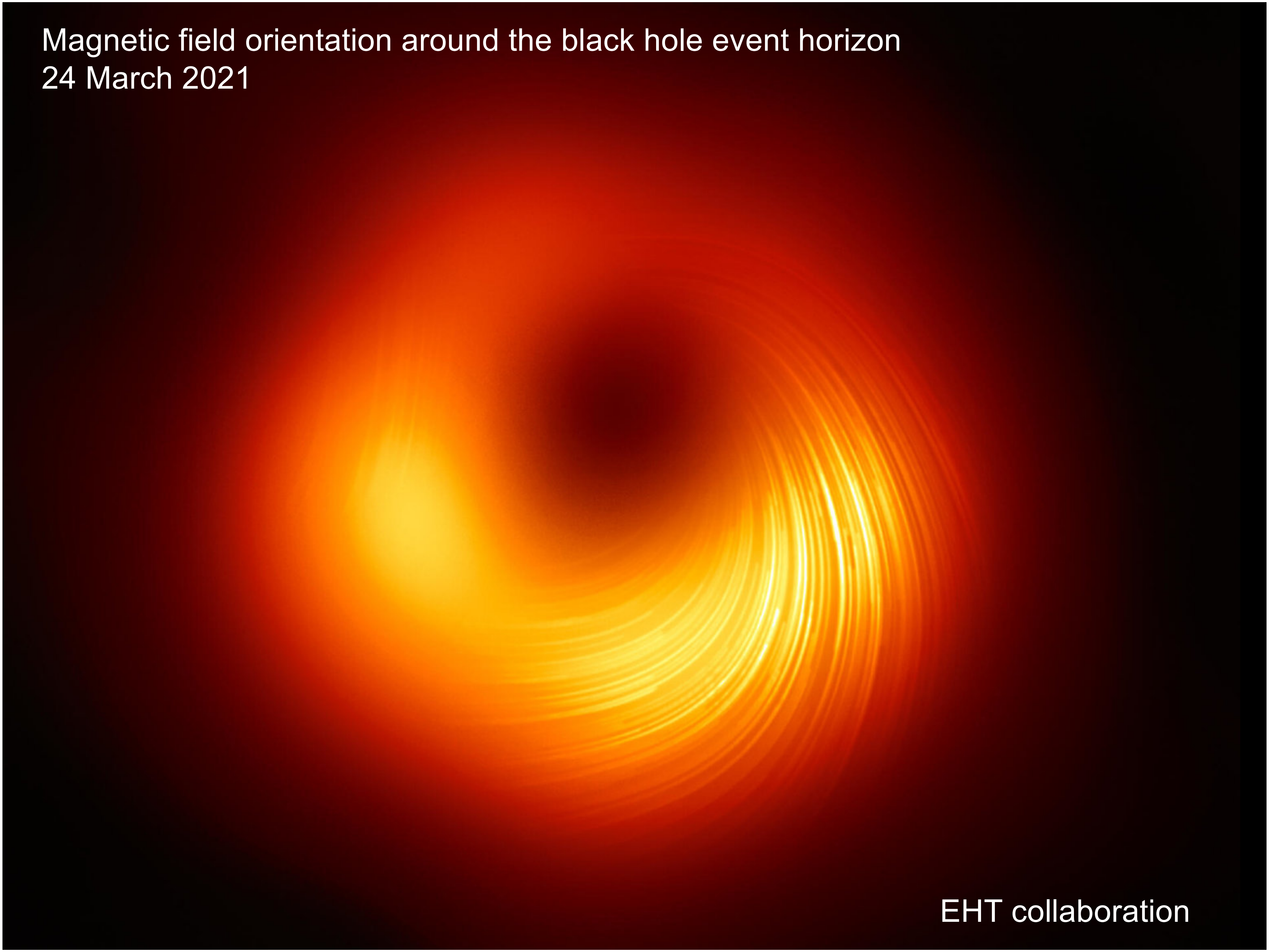


Eventhorizontelescope.org

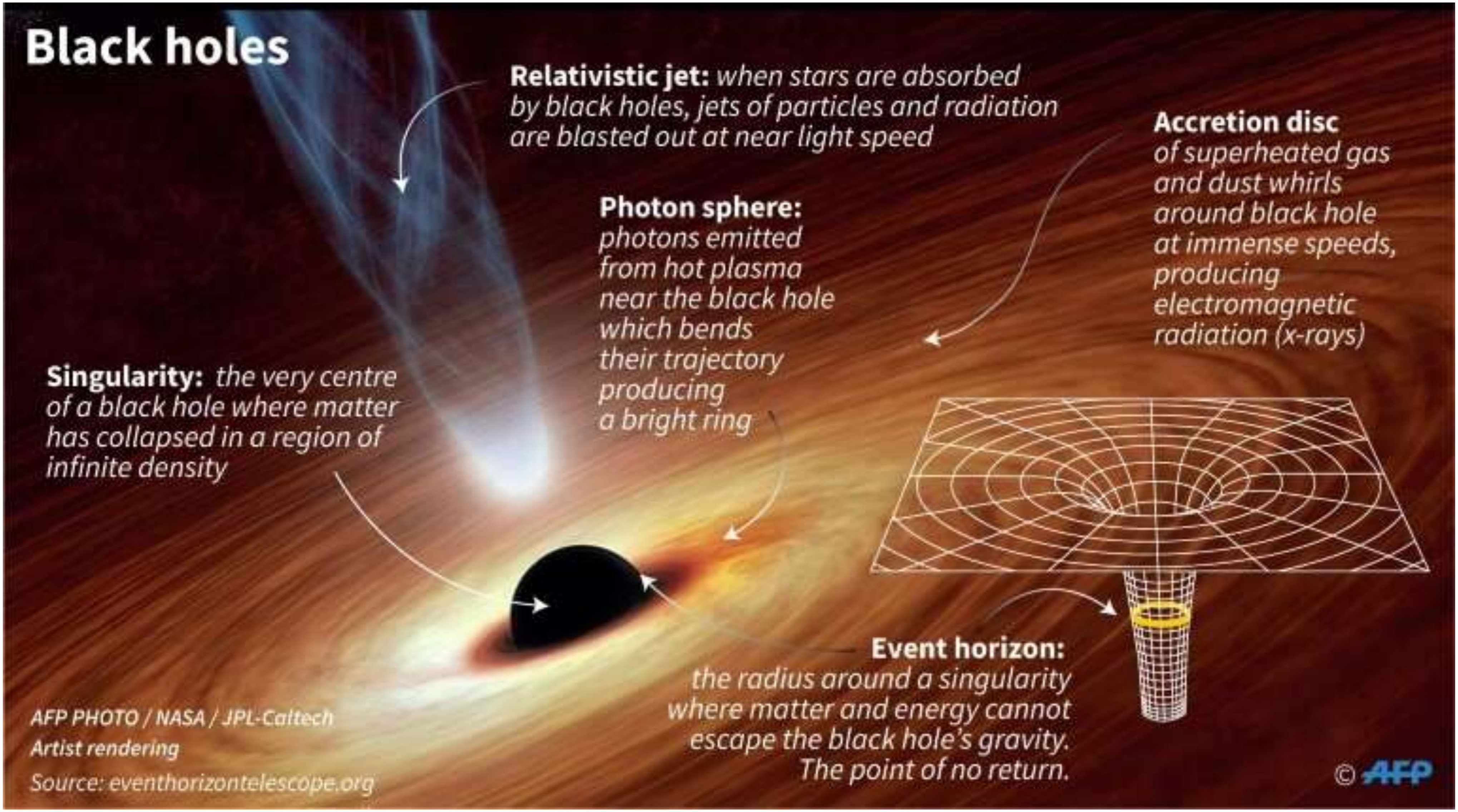
Radius of the event horizon:

$R_s = \frac{2GM}{c^2}$ M: mass of black hole

(Schwarzschild radius)

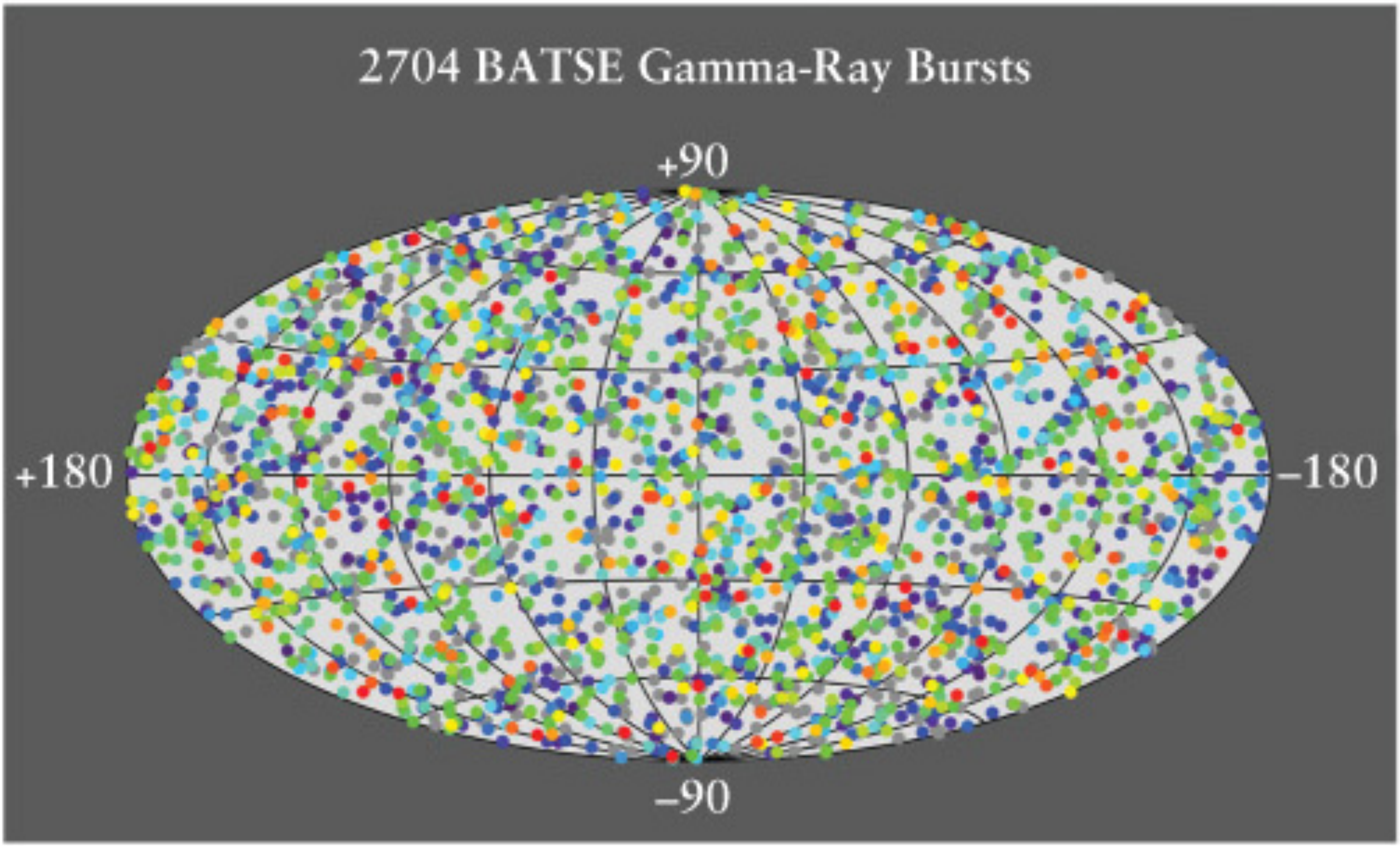

Magnetic field orientation around the black hole event horizon
24 March 2021
EHT collaboration

Wikipedia

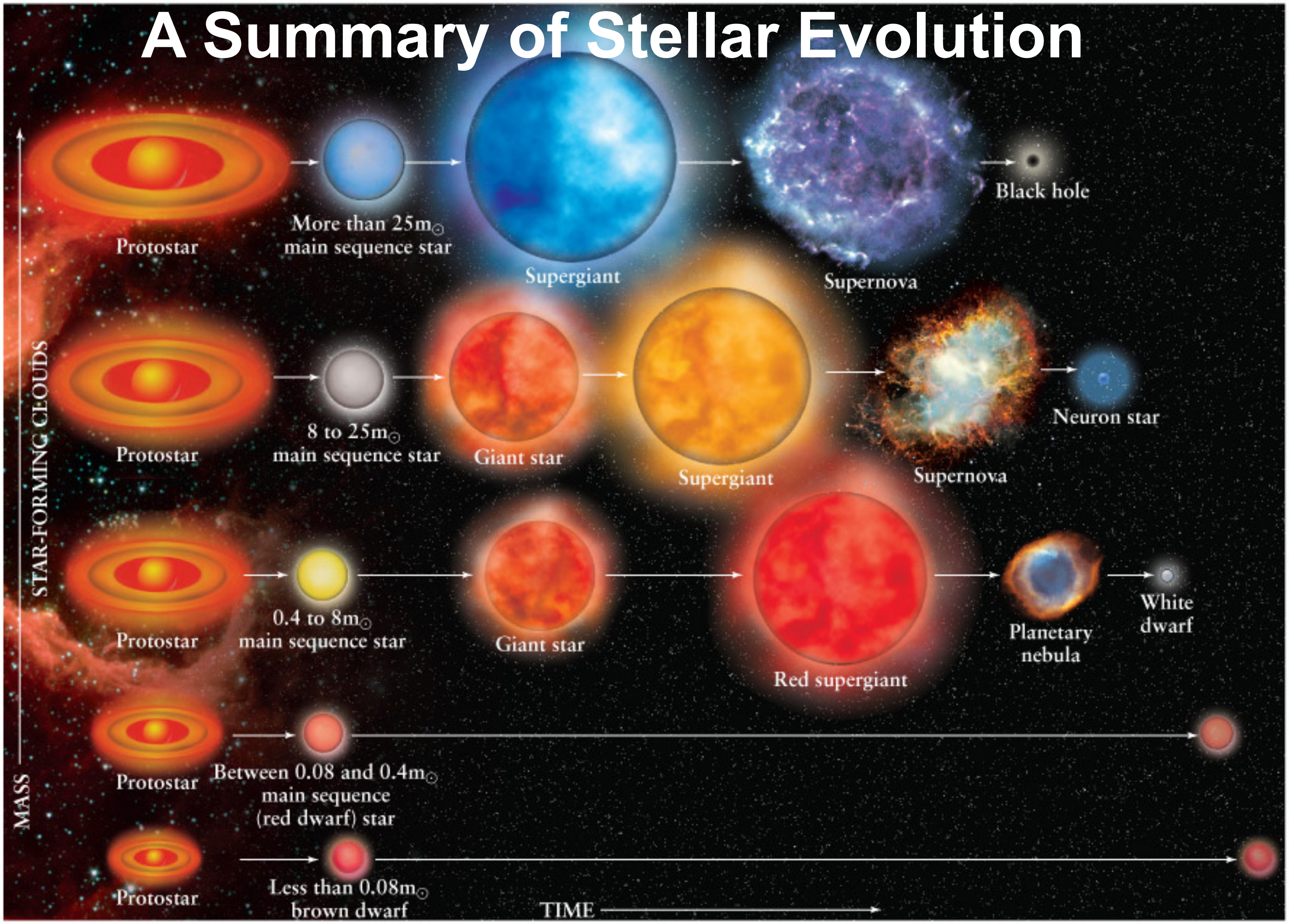


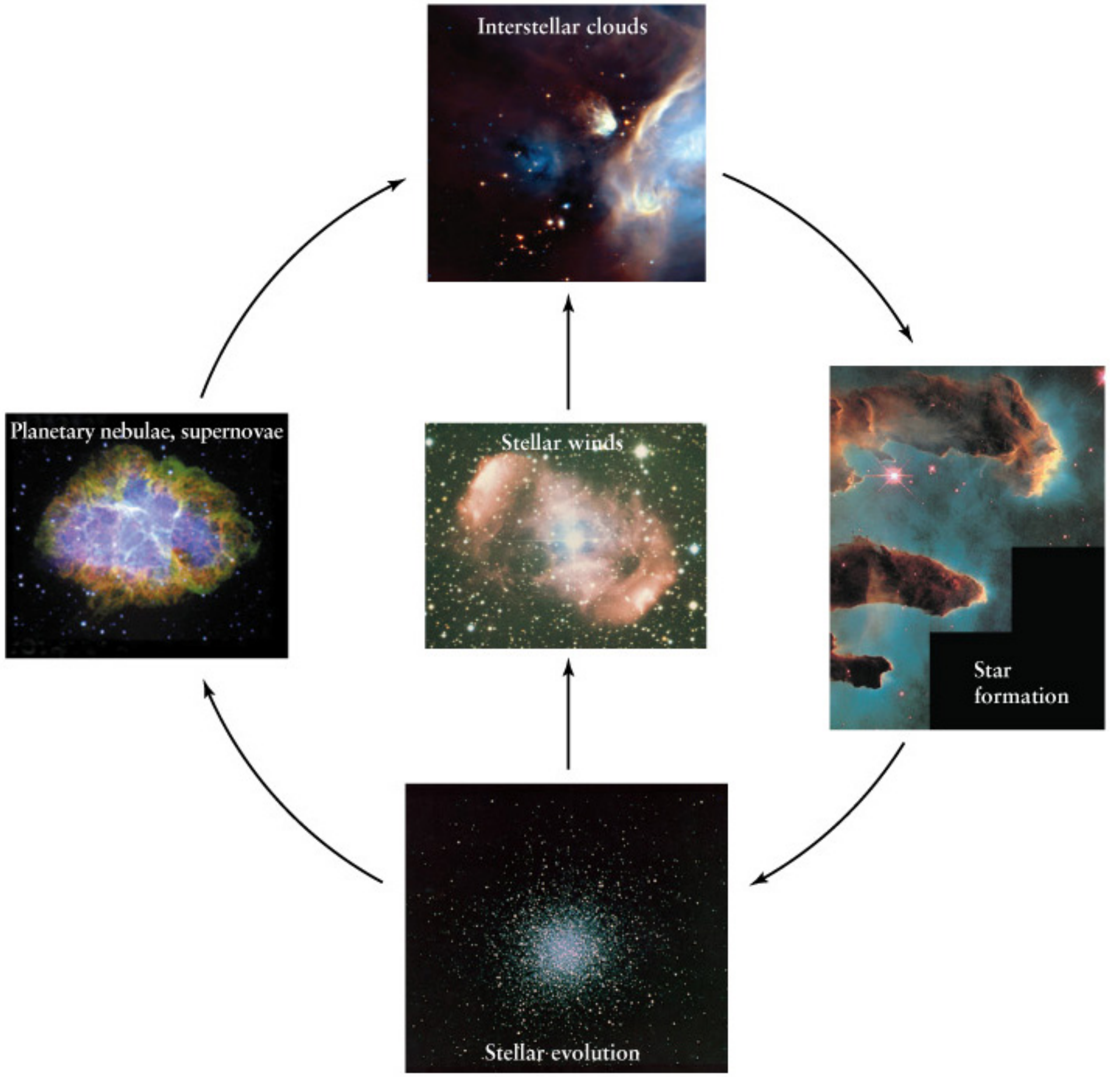
Eventhorizontelescope.org

The Most Powerful Known Gamma Ray Bursts



A Summary of Stellar Evolution




Interstellar clouds
Star
formation
Stellar evolution
Planetary nebulae, supernovae
Stellar winds

b